

Through The Looking Glass: Glass:



Faraday Conversion
in
Turbulent Blazar Jets



Nicholas MacDonald

Boston University


Talk Outline

A brief introduction to Faraday Conversion.

Can circular polarization be produced within a turbulent jet?

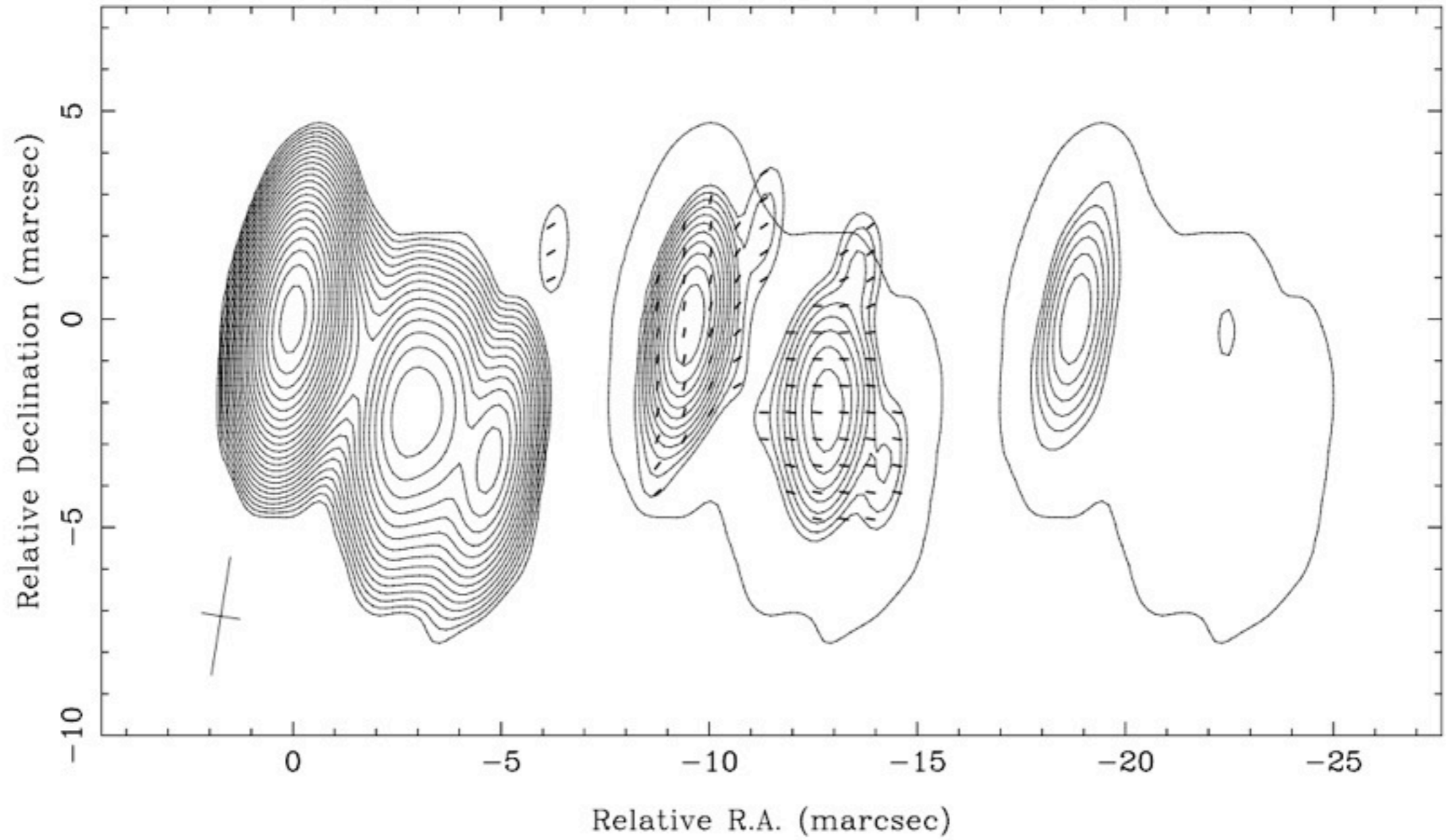
Malaga, Spain

June 2016



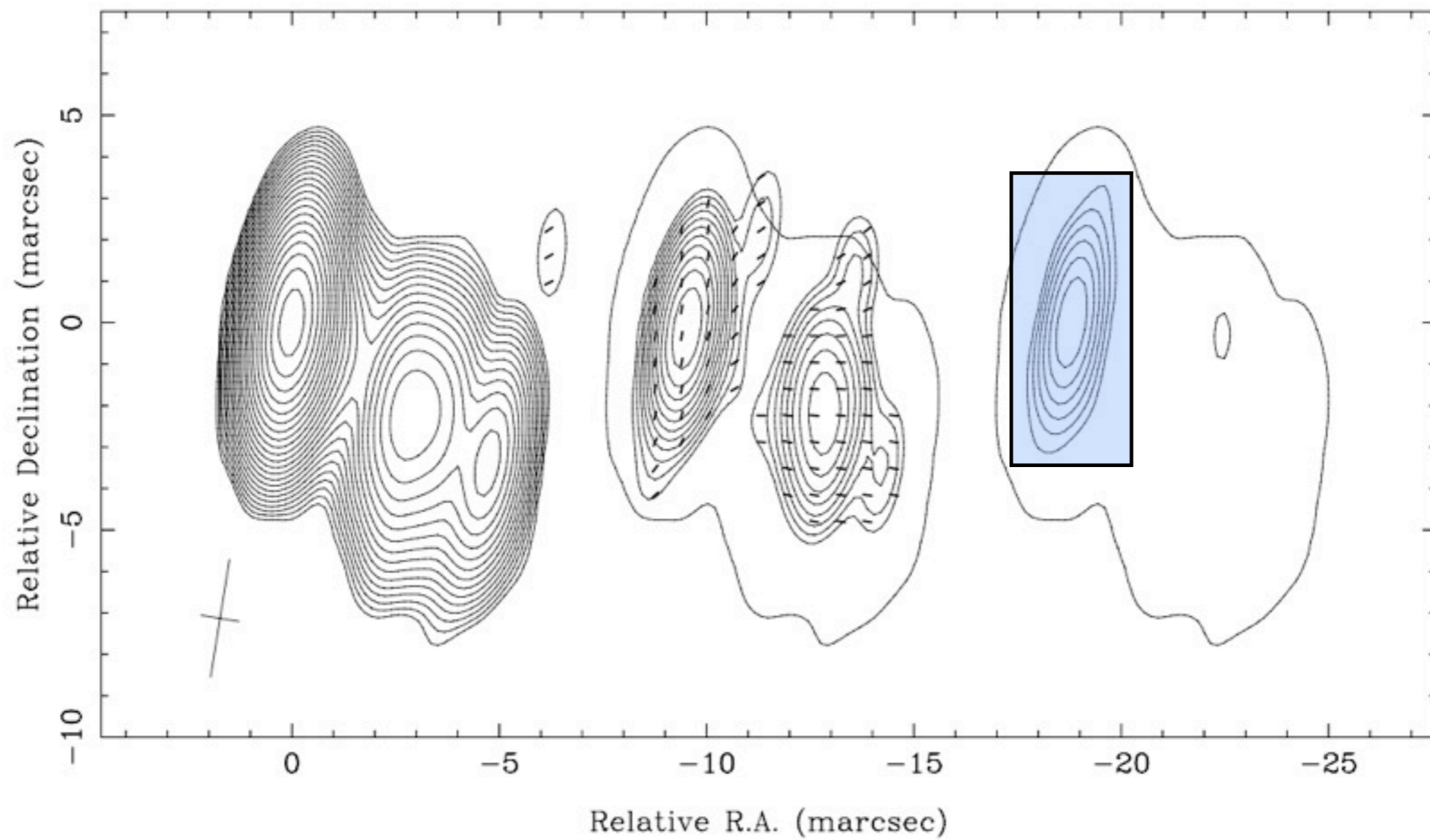
Birefringence

3C 279



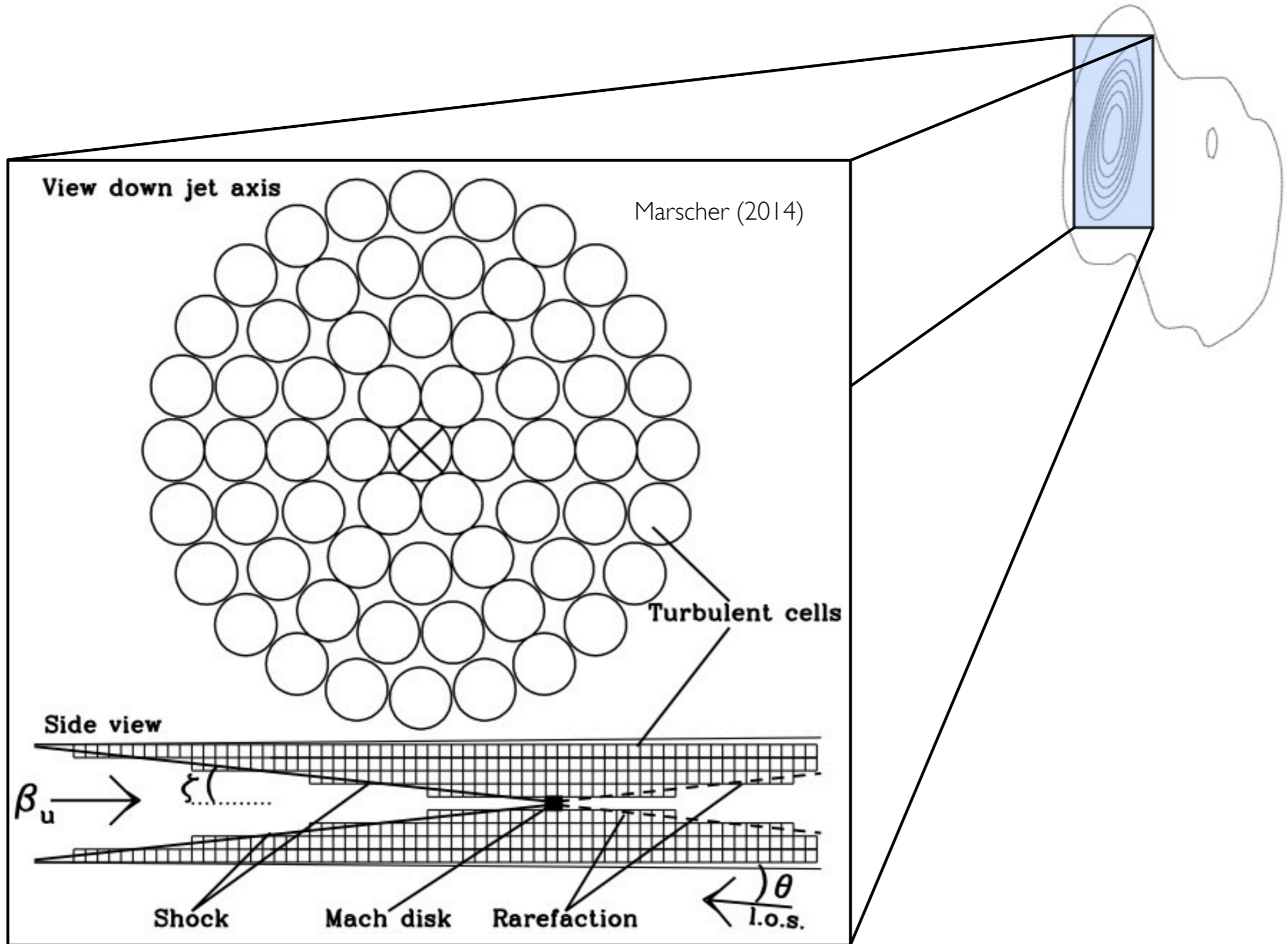
Homan et al. (2009)

3C 279

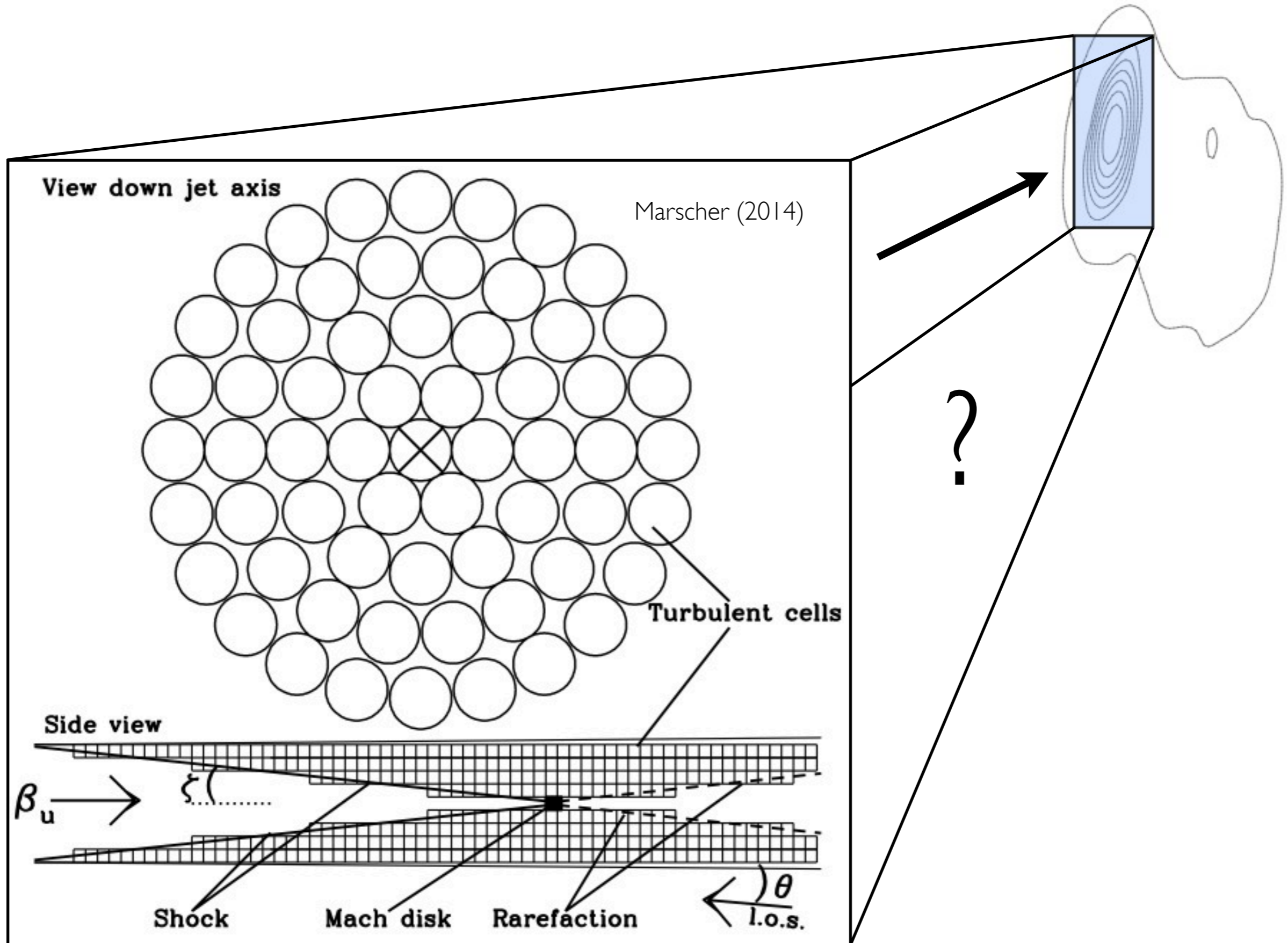


Homan et al. (2009)

The turbulent extreme multi-zone (TEMZ) code:



The turbulent extreme multi-zone (TEMZ) code:



1977

1977



6

1977



6

TRANSFER OF POLARIZED RADIATION IN SELF-ABSORBED SYNCHROTRON SOURCES.
I. RESULTS FOR A HOMOGENEOUS SOURCE

T. W. JONES

National Radio Astronomy Observatory*

AND

S. L. O'DELL

Department of Physics, University of California, San Diego

Received 1975 June 16; revised 1976 October 26

ABSTRACT

The solution to the equation of transfer of polarized radiation in a stationary, homogeneous, rarefied medium is applied to self-absorbed synchrotron sources. Relativistic electrons (independent of the presence of any cold plasma) can quite easily produce in such sources significant Faraday rotation and/or conversion of linear to circular polarization. Structural inhomogeneities do not obviate the importance of these phenomena in cosmic, compact nonthermal sources. Contrary to the calculation of Pacholczyk and Swihart, the circular polarization for a homogeneous source changes sign just below the self-absorption turnover as the source becomes opaque, even when polarization conversion dominates; however, for a physically realistic source, structural inhomogeneity may alter this behavior. The observational evidence bearing upon these effects is reviewed.

Subject headings: polarization — radiative transfer — radio sources: variable —
synchrotron radiation

1977



TRANSFER OF POLARIZED RADIATION IN SELF-ABSORBED SYNCHROTRON SOURCES.
I. RESULTS FOR A HOMOGENEOUS SOURCE

T. W. JONES

National Radio Astronomy Observatory*

AND

S. L. O'DELL

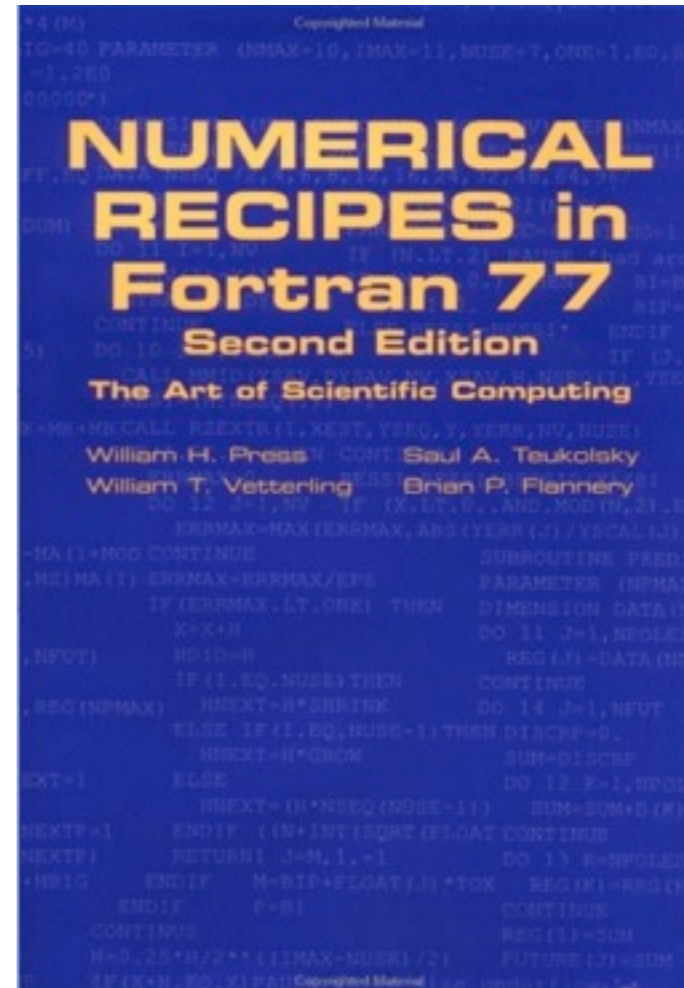
Department of Physics, University of California, San Diego

Received 1975 June 16; revised 1976 October 26

ABSTRACT

The solution to the equation of transfer of polarized radiation in a stationary, homogeneous, rarefied medium is applied to self-absorbed synchrotron sources. Relativistic electrons (independent of the presence of any cold plasma) can quite easily produce in such sources significant Faraday rotation and/or conversion of linear to circular polarization. Structural inhomogeneities do not obviate the importance of these phenomena in cosmic, compact nonthermal sources. Contrary to the calculation of Pacholczyk and Swihart, the circular polarization for a homogeneous source changes sign just below the self-absorption turnover as the source becomes opaque, even when polarization conversion dominates; however, for a physically realistic source, structural inhomogeneity may alter this behavior. The observational evidence bearing upon these effects is reviewed.

Subject headings: polarization — radiative transfer — radio sources: variable — synchrotron radiation



1977



TRANSFER OF POLARIZED RADIATION IN SELF-ABSORBED SYNCHROTRON SOURCES.
I. RESULTS FOR A HOMOGENEOUS SOURCE

T. W. JONES

National Radio Astronomy Observatory*

AND

S. L. O'DELL

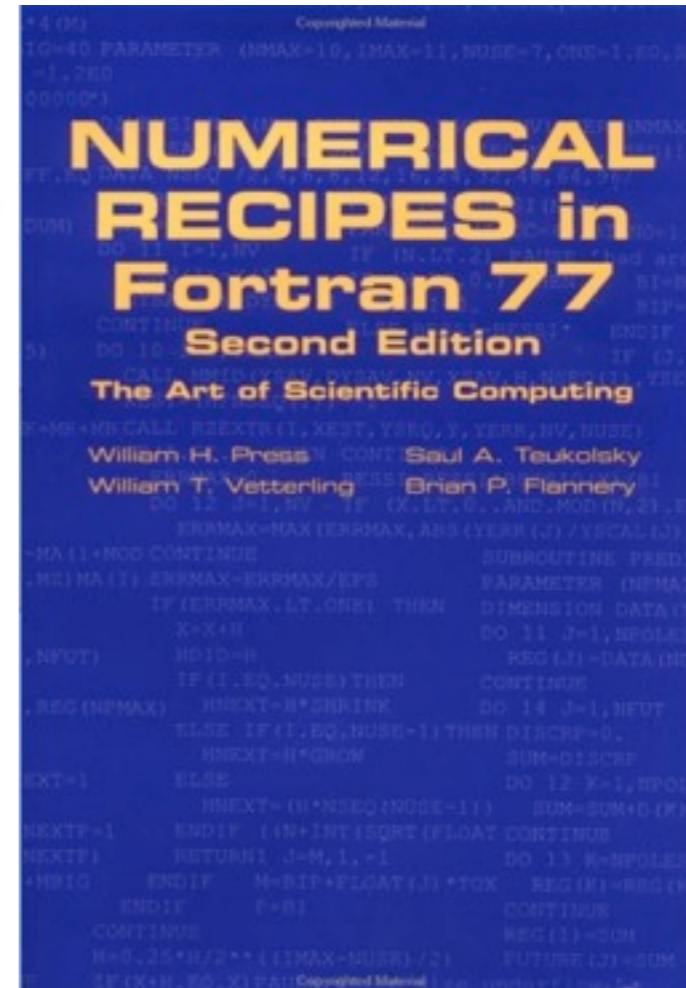
Department of Physics, University of California, San Diego

Received 1975 June 16; revised 1976 October 26

ABSTRACT

The solution to the equation of transfer of polarized radiation in a stationary, homogeneous, rarefied medium is applied to self-absorbed synchrotron sources. Relativistic electrons (independent of the presence of any cold plasma) can quite easily produce in such sources significant Faraday rotation and/or conversion of linear to circular polarization. Structural inhomogeneities do not obviate the importance of these phenomena in cosmic, compact nonthermal sources. Contrary to the calculation of Pacholczyk and Swihart, the circular polarization for a homogeneous source changes sign just below the self-absorption turnover as the source becomes opaque, even when polarization conversion dominates; however, for a physically realistic source, structural inhomogeneity may alter this behavior. The observational evidence bearing upon these effects is reviewed.

Subject headings: polarization — radiative transfer — radio sources: variable — synchrotron radiation



1977



Solving the Full Stokes Equations of Radiative Transfer

Stokes Parameters

$$\begin{array}{c}
 \left(\frac{d}{dl} + \kappa_I \right) \\
 \kappa_Q \\
 \kappa_U \\
 \kappa_V
 \end{array}
 \begin{array}{c}
 \kappa_Q \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_V^* \\
 \kappa_U^*
 \end{array}
 \begin{array}{c}
 \kappa_U \\
 \kappa_V^* \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_Q^*
 \end{array}
 \begin{array}{c}
 \kappa_V \\
 -\kappa_U^* \\
 \kappa_Q^* \\
 \left(\frac{d}{dl} + \kappa_I \right)
 \end{array}
 \begin{array}{c}
 I_\nu \\
 Q_\nu \\
 U_\nu \\
 V_\nu
 \end{array}
 =
 \begin{array}{c}
 \eta_\nu^I \\
 \eta_\nu^Q \\
 \eta_\nu^U \\
 \eta_\nu^V
 \end{array}$$

Jones & O'Dell (1977)

Emission Coefficients

$$\begin{array}{c}
 \left(\frac{d}{dl} + \kappa_I \right) \\
 \kappa_Q \\
 \kappa_U \\
 \kappa_V
 \end{array}
 \begin{array}{c}
 \kappa_Q \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_V^* \\
 \kappa_U^*
 \end{array}
 \begin{array}{c}
 \kappa_U \\
 \kappa_V^* \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_Q^*
 \end{array}
 \begin{array}{c}
 \kappa_V \\
 -\kappa_U^* \\
 \kappa_Q^* \\
 \left(\frac{d}{dl} + \kappa_I \right)
 \end{array}
 \begin{array}{c}
 I_\nu \\
 Q_\nu \\
 U_\nu \\
 V_\nu
 \end{array}
 =
 \begin{array}{c}
 \eta_\nu^I \\
 \eta_\nu^Q \\
 \eta_\nu^U \\
 \eta_\nu^V
 \end{array}$$

Jones & O'Dell (1977)

Absorption Coefficients

$$\left(\frac{d}{dl} + \boxed{\kappa_I} \right)$$

$$\begin{array}{|c|} \hline \kappa_Q \\ \hline \kappa_U \\ \hline \kappa_V \\ \hline \end{array}$$

$$\begin{array}{|c|c|c|} \hline \kappa_Q & \kappa_U & \kappa_V \\ \hline \end{array}$$

$$\left(\frac{d}{dl} + \boxed{\kappa_I} \right)$$

$$-\kappa_V^*$$

$$\kappa_U^*$$

$$\left(\frac{d}{dl} + \boxed{\kappa_I} \right)$$

$$-\kappa_Q^*$$

$$\left(\frac{d}{dl} + \boxed{\kappa_I} \right)$$

$$\kappa_V^*$$

$$-\kappa_U^*$$

$$\kappa_Q^*$$

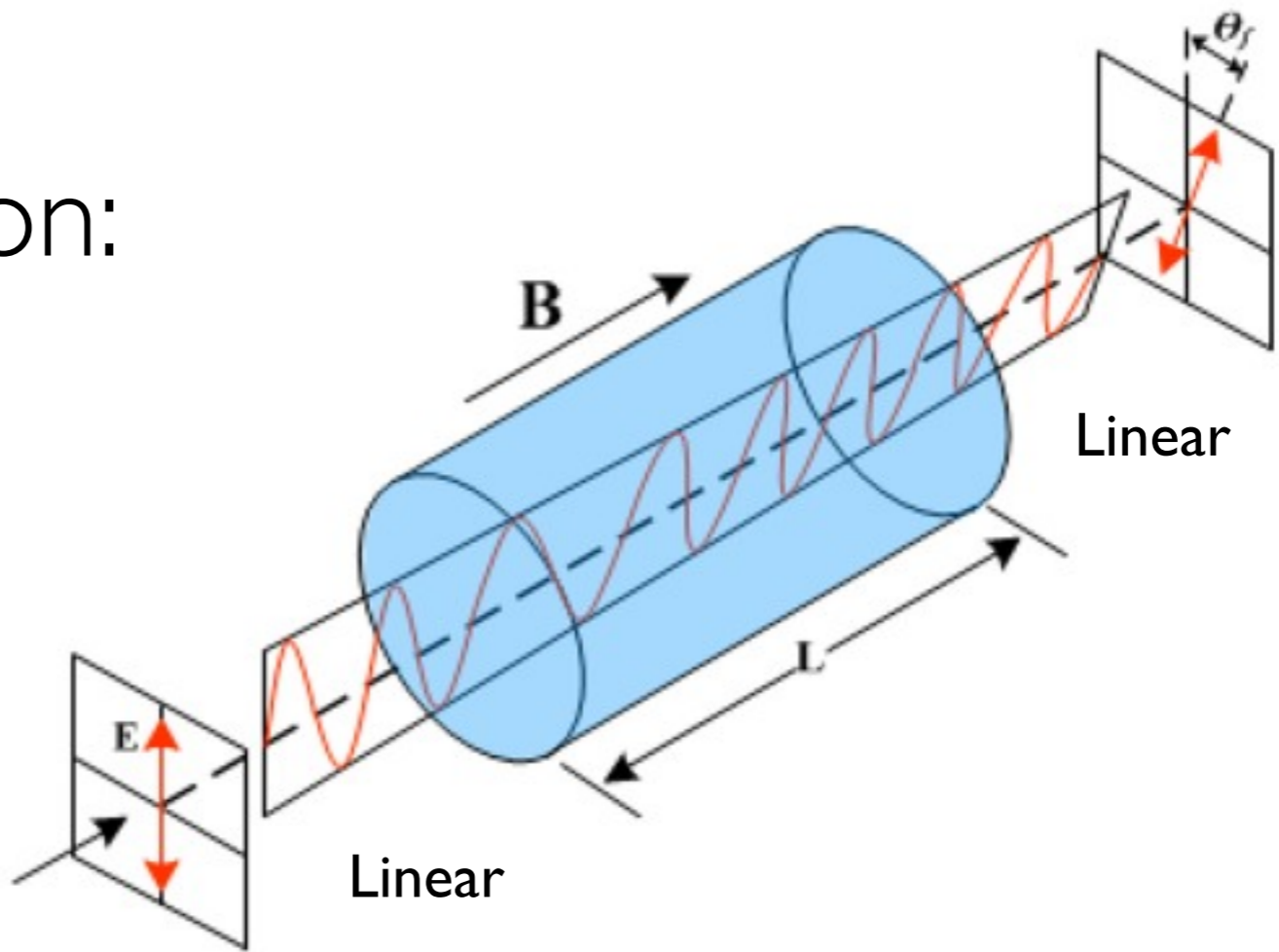
$$\begin{array}{|c|} \hline I_\nu \\ \hline Q_\nu \\ \hline U_\nu \\ \hline V_\nu \\ \hline \end{array} = \begin{array}{|c|} \hline \eta_\nu^I \\ \hline \eta_\nu^Q \\ \hline \eta_\nu^U \\ \hline \eta_\nu^V \\ \hline \end{array}$$

Jones & O'Dell (1977)

$$\begin{array}{c}
 \left(\frac{d}{dl} + \kappa_I \right) \\
 \kappa_Q \\
 \kappa_U \\
 \kappa_V
 \end{array}
 \begin{array}{c}
 \kappa_Q \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 \boxed{-\kappa_V^*} \\
 \kappa_U^*
 \end{array}
 \begin{array}{c}
 \kappa_U \\
 \boxed{\kappa_V^*} \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_Q^*
 \end{array}
 \begin{array}{c}
 \kappa_V \\
 -\kappa_U^* \\
 \kappa_Q^* \\
 \left(\frac{d}{dl} + \kappa_I \right)
 \end{array}
 \left| \begin{array}{c}
 I_\nu \\
 Q_\nu \\
 U_\nu \\
 V_\nu
 \end{array} \right| = \left| \begin{array}{c}
 \eta_\nu^I \\
 \eta_\nu^Q \\
 \eta_\nu^U \\
 \eta_\nu^V
 \end{array} \right|$$

Jones & O'Dell (1977)

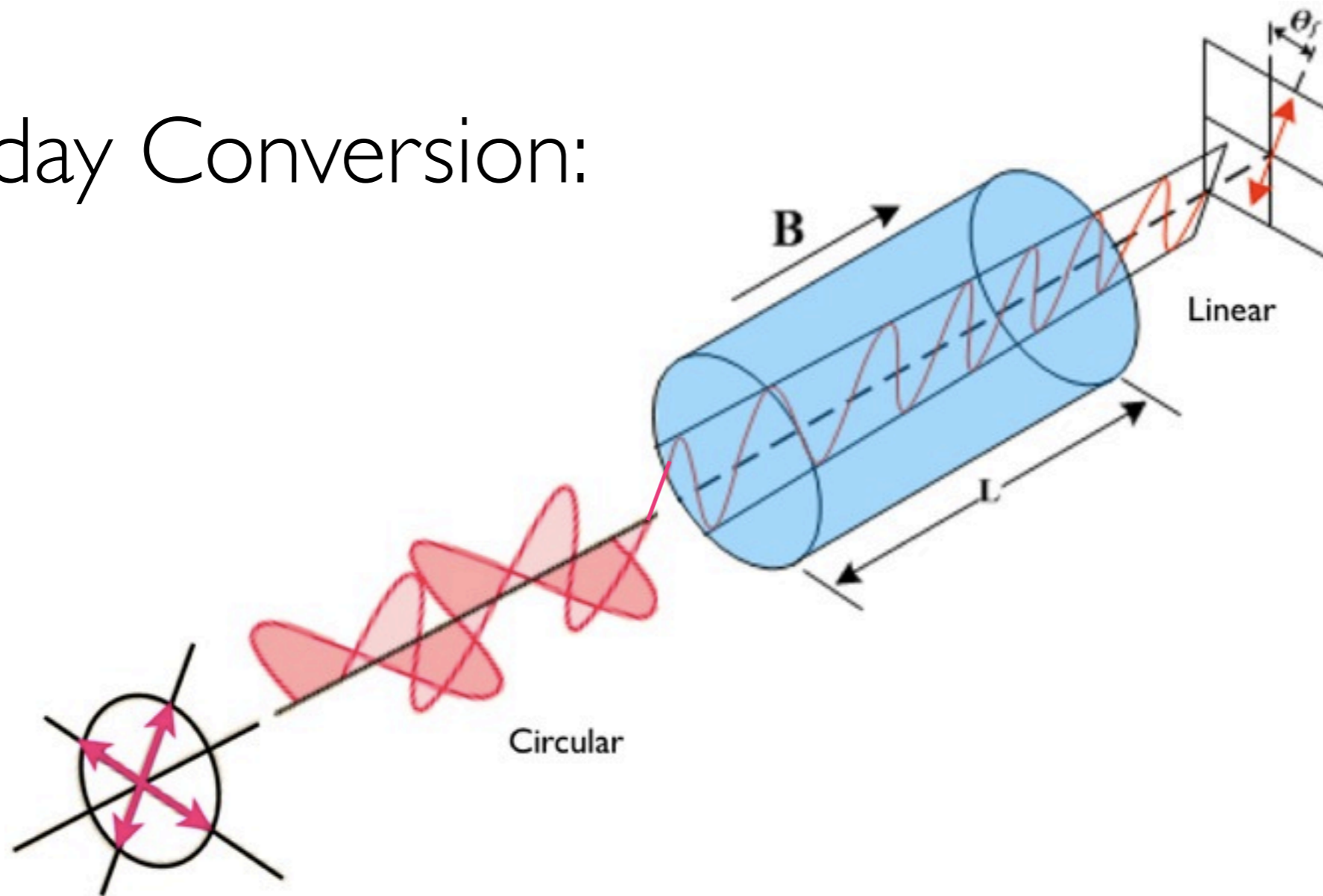
Faraday Rotation:



$$\begin{array}{c}
 \left(\frac{d}{dl} + \kappa_I \right) \\
 \kappa_Q \\
 \kappa_U \\
 \kappa_V
 \end{array}
 \begin{array}{c}
 \kappa_Q \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_V^* \\
 \boxed{\kappa_U^* \quad -\kappa_Q^*}
 \end{array}
 \begin{array}{c}
 \kappa_U \\
 \kappa_V^* \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 \left(\frac{d}{dl} + \kappa_I \right)
 \end{array}
 \begin{array}{c}
 \kappa_V \\
 \boxed{-\kappa_U^* \\ \kappa_Q^*} \\
 \left(\frac{d}{dl} + \kappa_I \right)
 \end{array}
 \begin{array}{c}
 I_\nu \\
 Q_\nu \\
 U_\nu \\
 V_\nu
 \end{array}
 =
 \begin{array}{c}
 \eta_\nu^I \\
 \eta_\nu^Q \\
 \eta_\nu^U \\
 \eta_\nu^V
 \end{array}$$

Jones & O'Dell (1977)

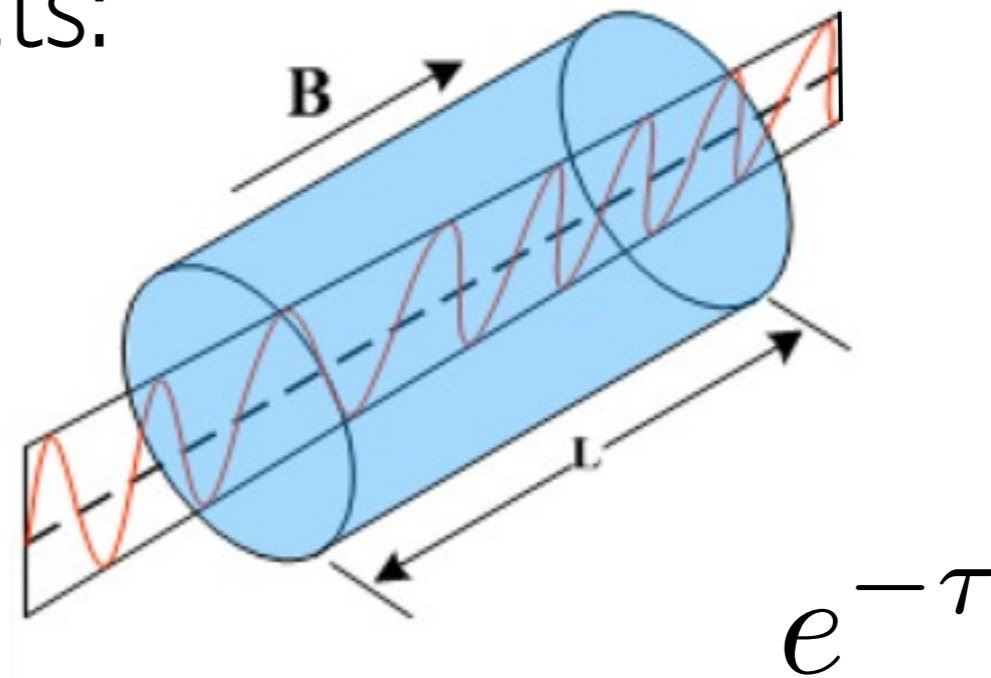
Faraday Conversion:



$$\begin{array}{c}
 \left(\frac{d}{dl} + \kappa_I \right) \\
 \kappa_Q \\
 \kappa_U \\
 \kappa_V
 \end{array}
 \begin{array}{c}
 \kappa_Q \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_V^* \\
 \kappa_U^*
 \end{array}
 \begin{array}{c}
 \kappa_U \\
 \kappa_V^* \\
 \left(\frac{d}{dl} + \kappa_I \right) \\
 -\kappa_Q^*
 \end{array}
 \begin{array}{c}
 \kappa_V \\
 -\kappa_U^* \\
 \kappa_Q^* \\
 \left(\frac{d}{dl} + \kappa_I \right)
 \end{array}
 \begin{array}{c}
 I_\nu \\
 Q_\nu \\
 U_\nu \\
 V_\nu
 \end{array}
 =
 \begin{array}{c}
 \eta_\nu^I \\
 \eta_\nu^Q \\
 \eta_\nu^U \\
 \eta_\nu^V
 \end{array}$$

Jones & O'Dell (1977)

Optical Depth Effects:



The Full Stokes Equations for Radiative Transfer of Polarized Emission:

$$\begin{array}{|c|} \hline I_\nu \\ \hline Q_\nu \\ \hline U_\nu \\ \hline V_\nu \\ \hline \end{array} = \begin{array}{|c|} \hline I_\nu^\infty \\ \hline Q_\nu^\infty \\ \hline U_\nu^\infty \\ \hline V_\nu^\infty \\ \hline \end{array} + e^{-\tau} \left\{ \begin{array}{l} \text{A lot of math!} \\ (B, \theta, L, \nu, n_o, \alpha, \gamma_{\min}) \end{array} \right\} \begin{array}{|c|} \hline I_\nu^0 - I_\nu^\infty \\ \hline Q_\nu^0 - Q_\nu^\infty \\ \hline U_\nu^0 - U_\nu^\infty \\ \hline V_\nu^0 - V_\nu^\infty \\ \hline \end{array}$$

Jones & O'Dell (1977)

The Full Stokes Equations for Radiative Transfer of Polarized Emission:

$$\begin{array}{|c|} \hline I_\nu \\ \hline Q_\nu \\ \hline U_\nu \\ \hline V_\nu \\ \hline \end{array} = \begin{array}{|c|} \hline I_\nu^\infty \\ \hline Q_\nu^\infty \\ \hline U_\nu^\infty \\ \hline V_\nu^\infty \\ \hline \end{array} + e^{-\tau} \left\{ \begin{array}{l} \text{A lot of math!} \\ (B, \theta, L, \nu, n_o, \alpha, \gamma_{\min}) \end{array} \right\} \begin{array}{|c|} \hline I_\nu^0 - I_\nu^\infty \\ \hline Q_\nu^0 - Q_\nu^\infty \\ \hline U_\nu^0 - U_\nu^\infty \\ \hline V_\nu^0 - V_\nu^\infty \\ \hline \end{array}$$

Jones & O'Dell (1977)

The Full Stokes Equations for Radiative Transfer of Polarized Emission:

$$\begin{pmatrix} I_\nu \\ Q_\nu \\ U_\nu \\ V_\nu \end{pmatrix} = \begin{pmatrix} I_\nu^\infty \\ Q_\nu^\infty \\ U_\nu^\infty \\ V_\nu^\infty \end{pmatrix} + e^{-\tau} \left\{ \begin{array}{l} \text{A lot of math!} \\ (B, \theta, L, \nu, n_o, \alpha, \gamma_{\min}) \end{array} \right\} \begin{pmatrix} I_\nu^0 - I_\nu^\infty \\ Q_\nu^0 - Q_\nu^\infty \\ U_\nu^0 - U_\nu^\infty \\ V_\nu^0 - V_\nu^\infty \end{pmatrix}$$

Jones & O'Dell (1977)

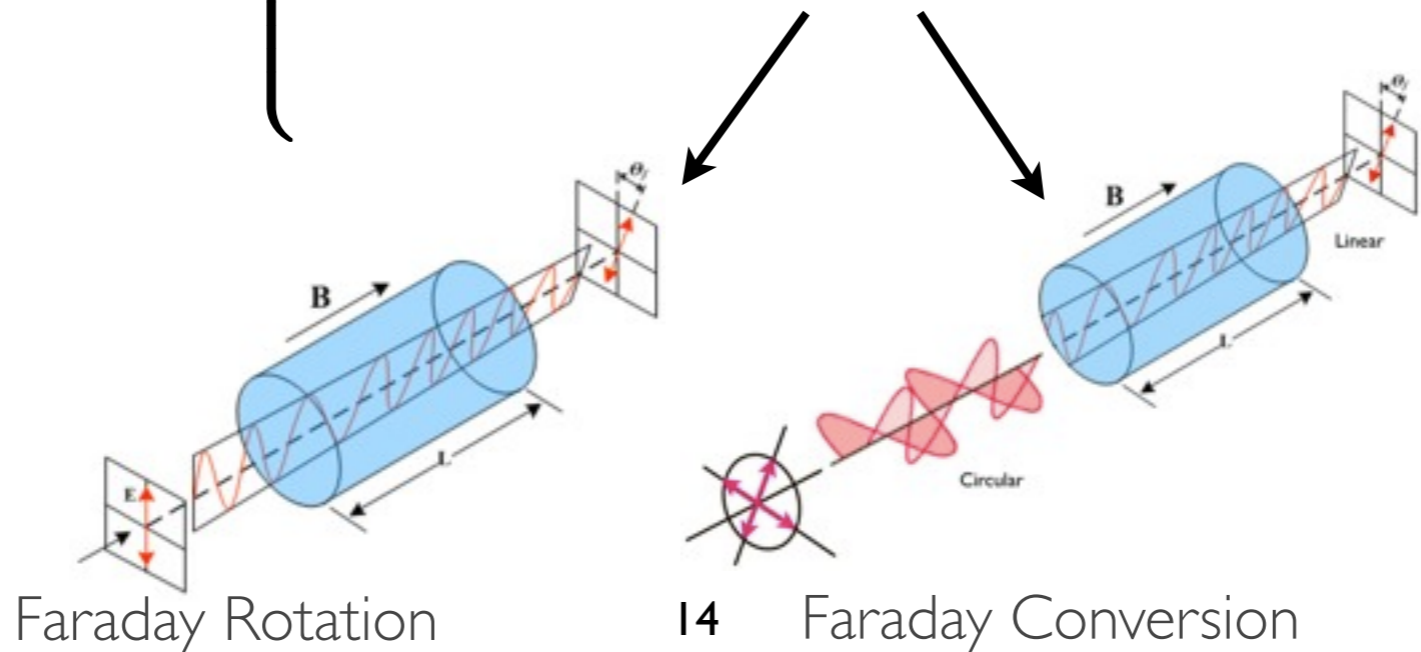
The Full Stokes Equations for Radiative Transfer of Polarized Emission:

$$\begin{pmatrix} I_\nu \\ Q_\nu \\ U_\nu \\ V_\nu \end{pmatrix} = \begin{pmatrix} I_\nu^\infty \\ Q_\nu^\infty \\ U_\nu^\infty \\ V_\nu^\infty \end{pmatrix} + e^{-\tau} \left\{ \begin{array}{c} \text{A lot of math!} \\ (B, \theta, L, \nu, n_o, \alpha, \gamma_{\min}) \end{array} \right\} \begin{pmatrix} I_\nu^0 - I_\nu^\infty \\ Q_\nu^0 - Q_\nu^\infty \\ U_\nu^0 - U_\nu^\infty \\ V_\nu^0 - V_\nu^\infty \end{pmatrix}$$

Jones & O'Dell (1977)

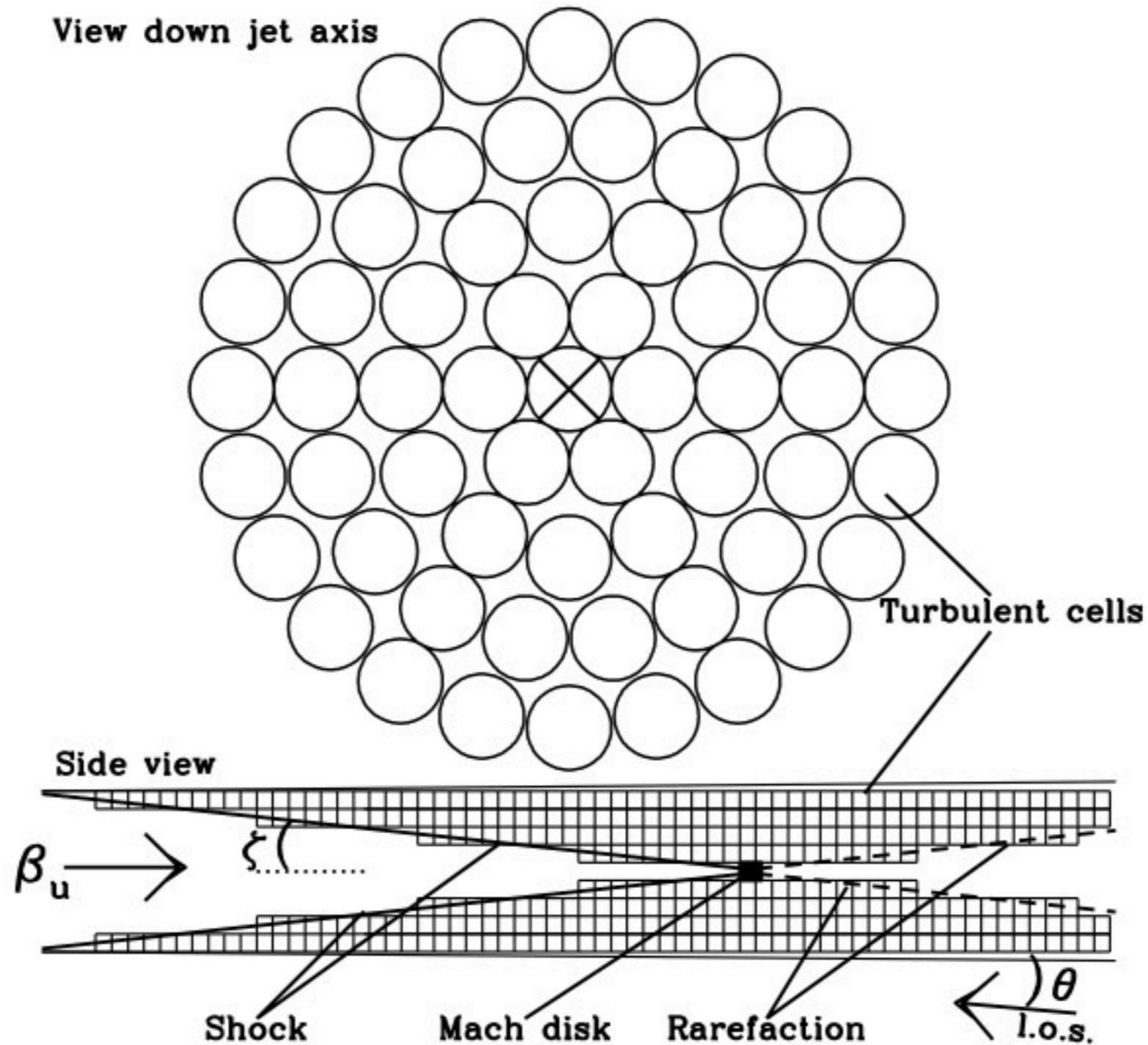
The Full Stokes Equations for Radiative Transfer of Polarized Emission:

$$\begin{pmatrix} I_\nu \\ Q_\nu \\ U_\nu \\ V_\nu \end{pmatrix} = \begin{pmatrix} I_\nu^\infty \\ Q_\nu^\infty \\ U_\nu^\infty \\ V_\nu^\infty \end{pmatrix} + e^{-\tau} \left\{ \begin{array}{c} \text{A lot of math!} \\ (B, \theta, L, \nu, n_o, \alpha, \gamma_{\min}) \end{array} \right\} \begin{pmatrix} I_\nu^0 - I_\nu^\infty \\ Q_\nu^0 - Q_\nu^\infty \\ U_\nu^0 - U_\nu^\infty \\ V_\nu^0 - V_\nu^\infty \end{pmatrix}$$



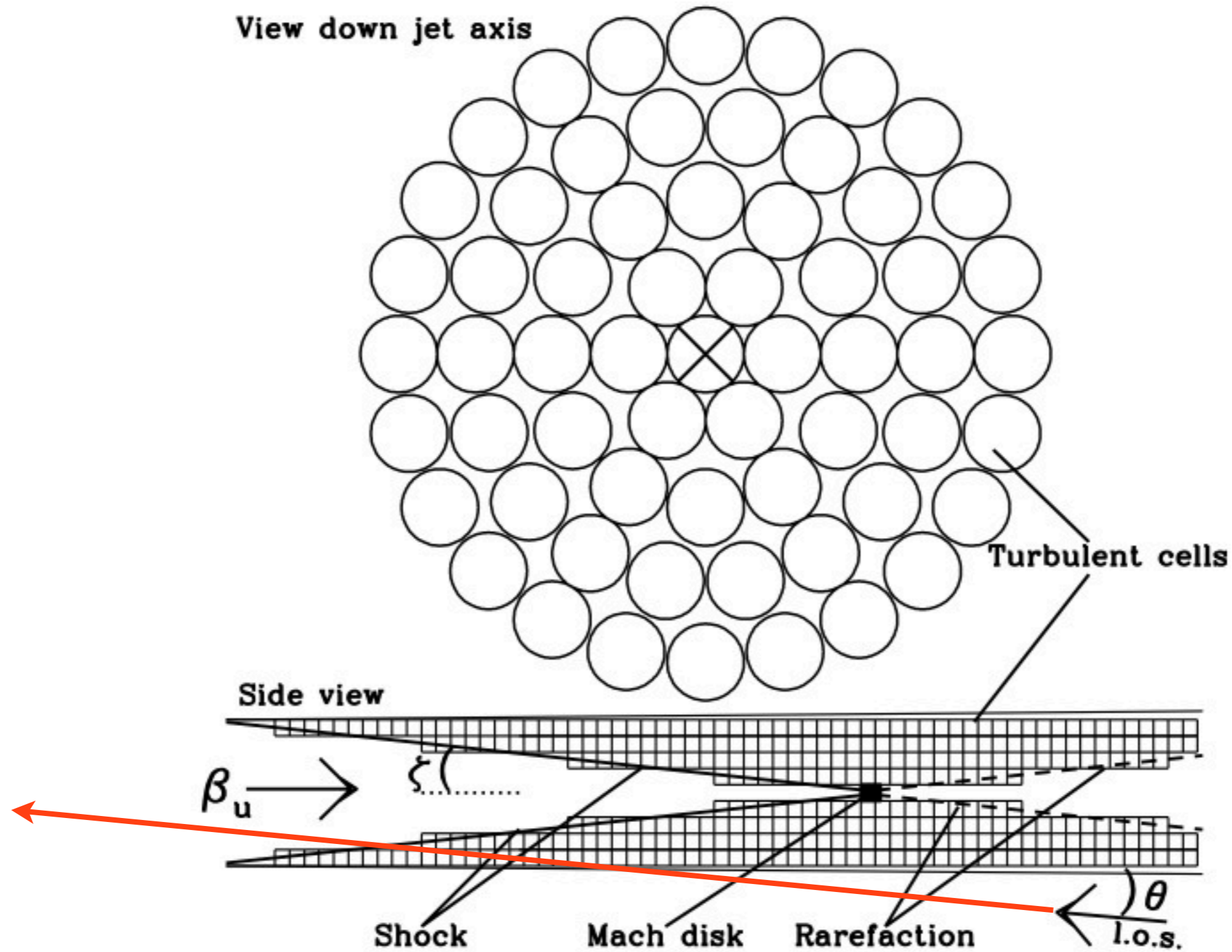
Jones & O'Dell (1977)

The turbulent extreme multi-zone (TEMZ) code:



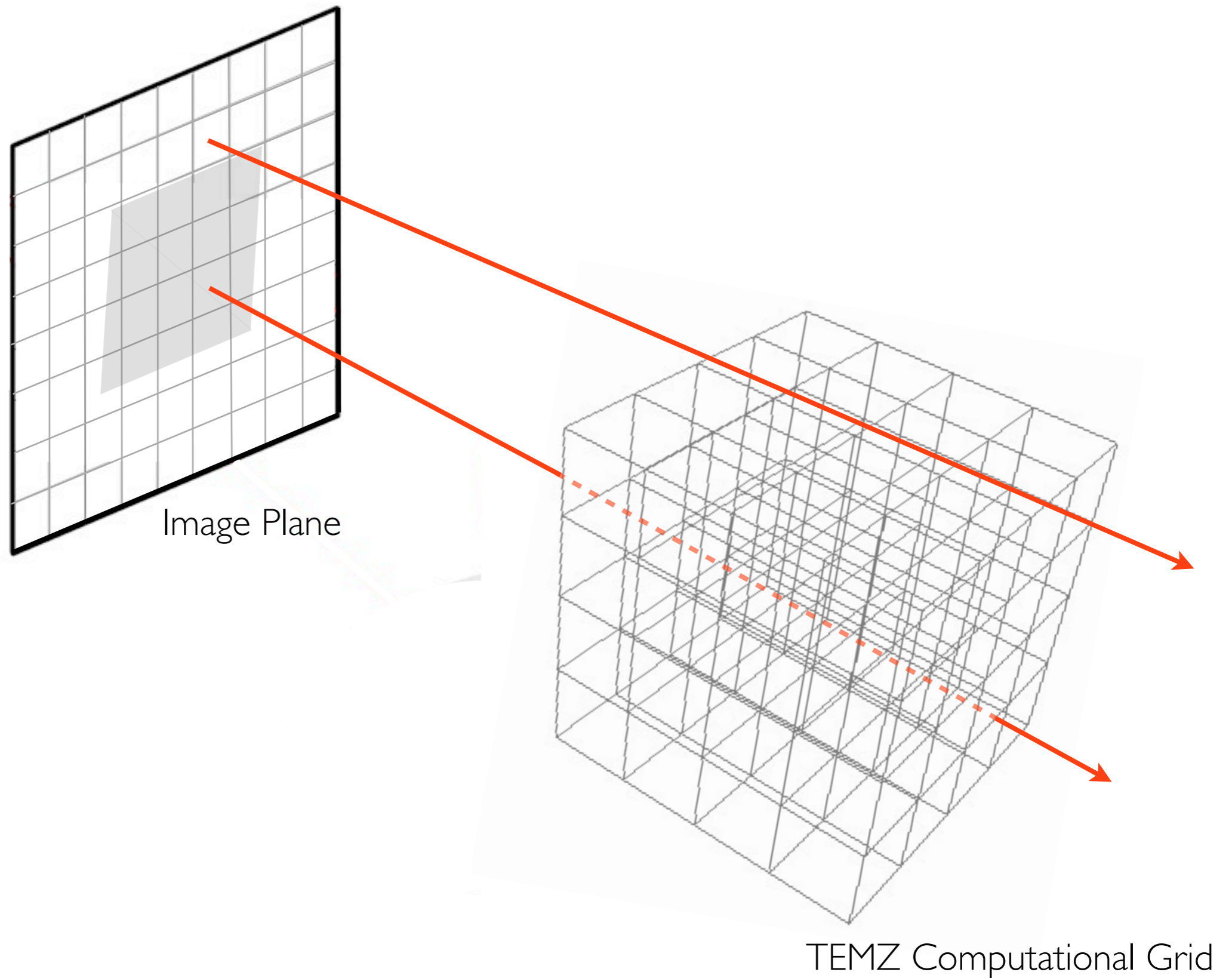
Marscher (2014)

The turbulent extreme multi-zone (TEMZ) code:

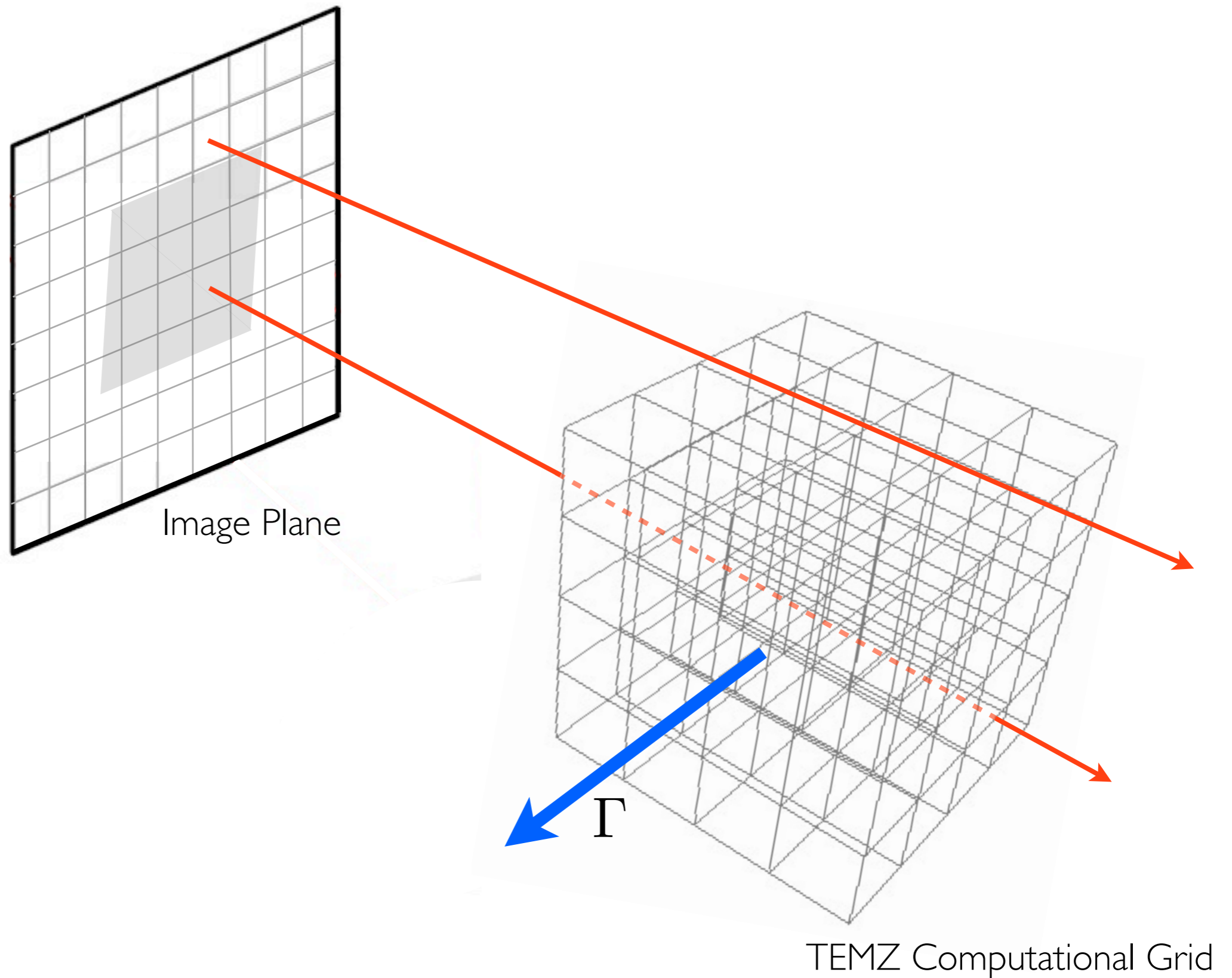


Marscher (2014)

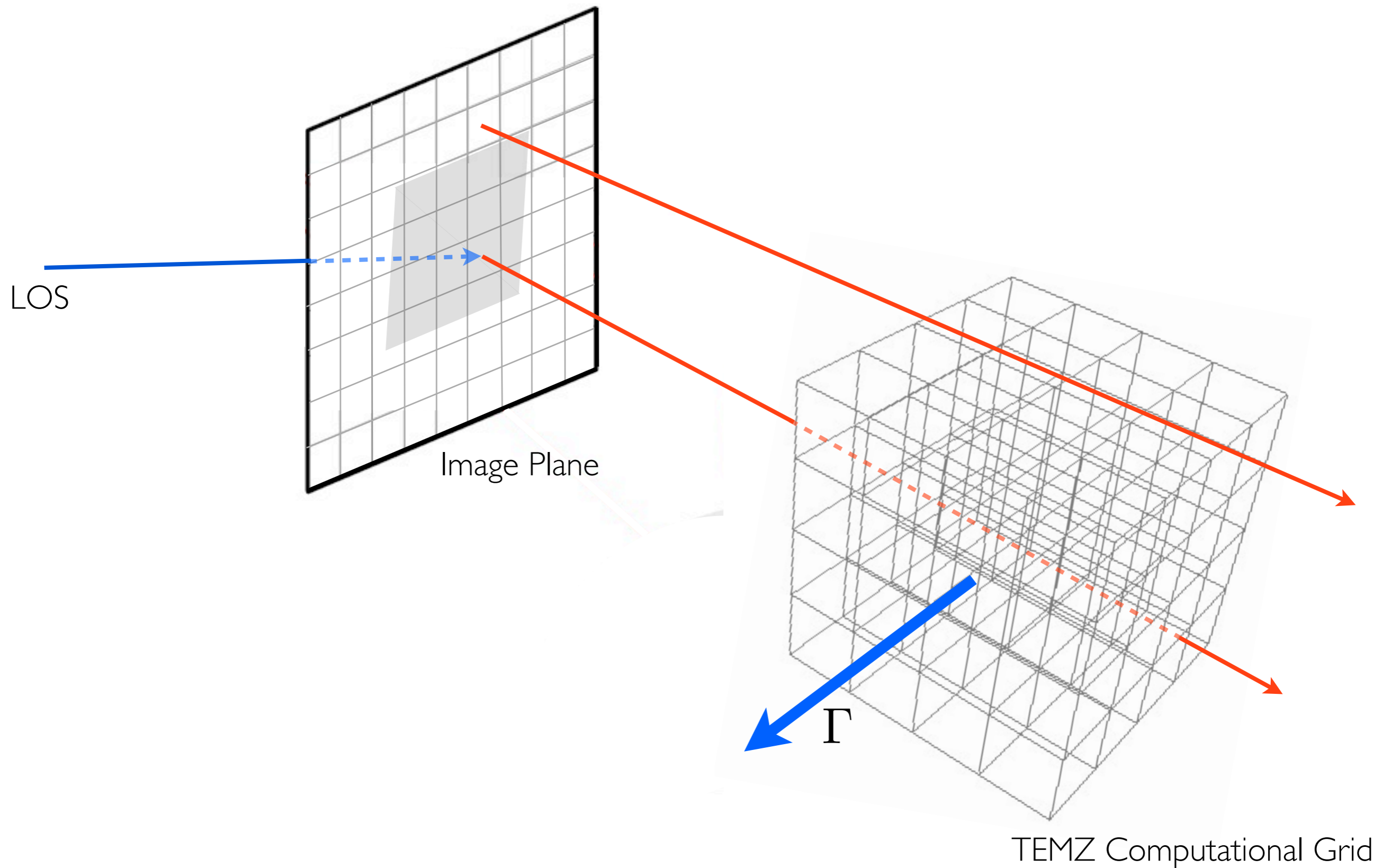
Ray Tracing Algorithm:



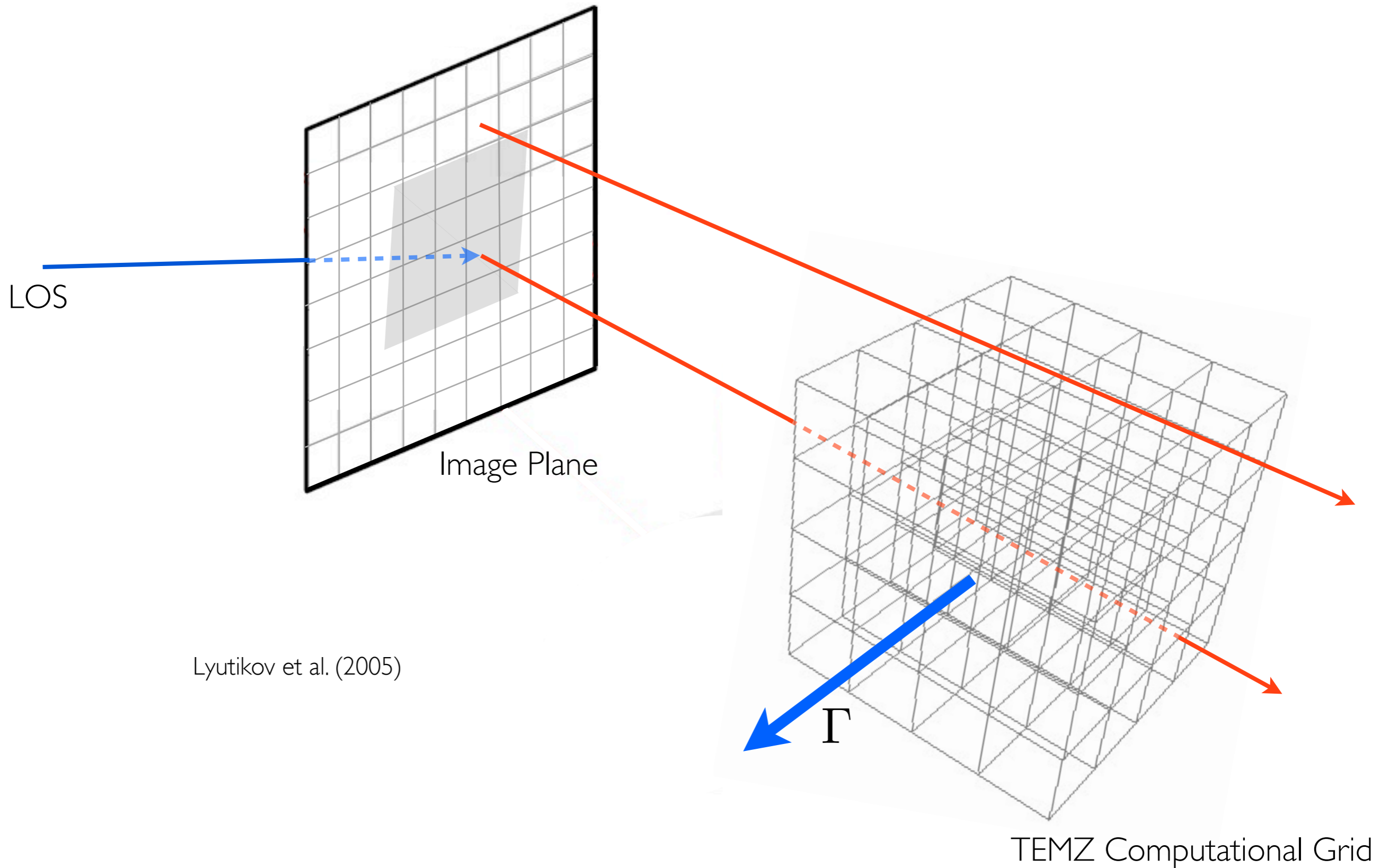
Ray Tracing Algorithm:



Ray Tracing Algorithm:



Ray Tracing Algorithm:

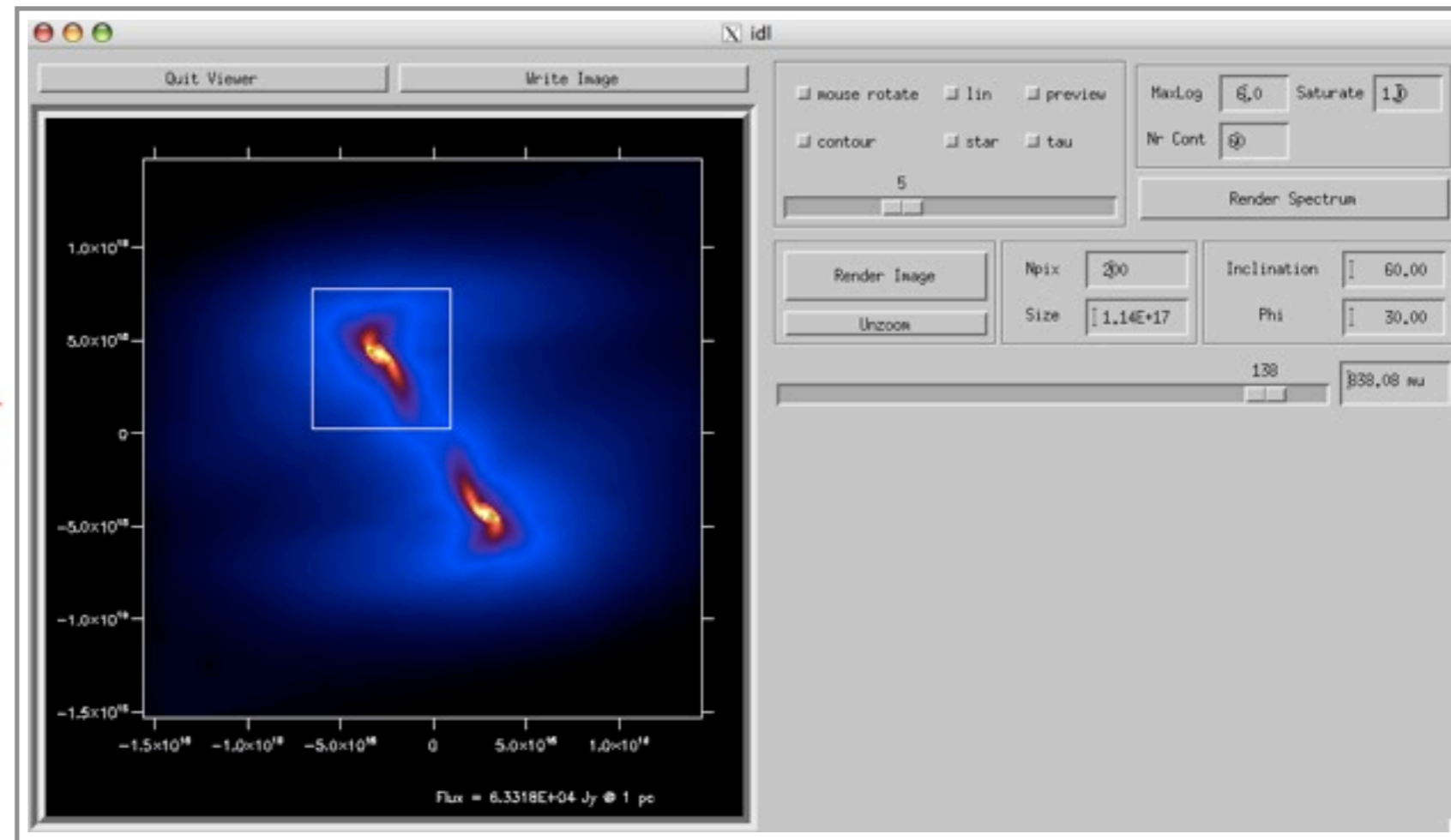
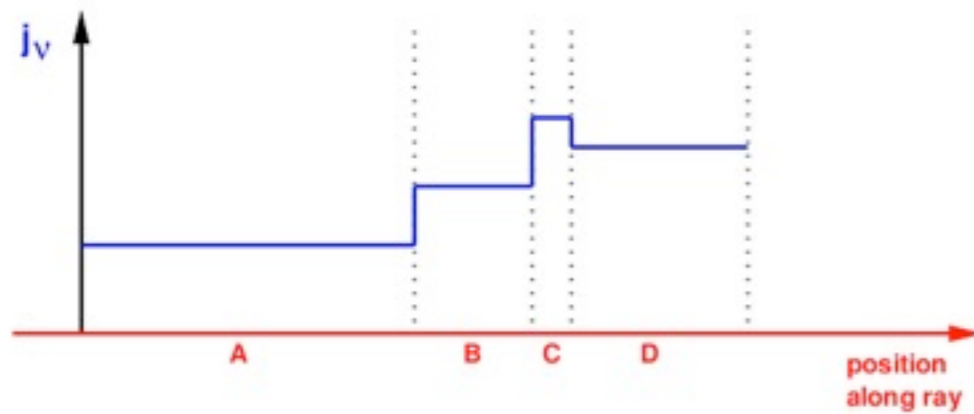
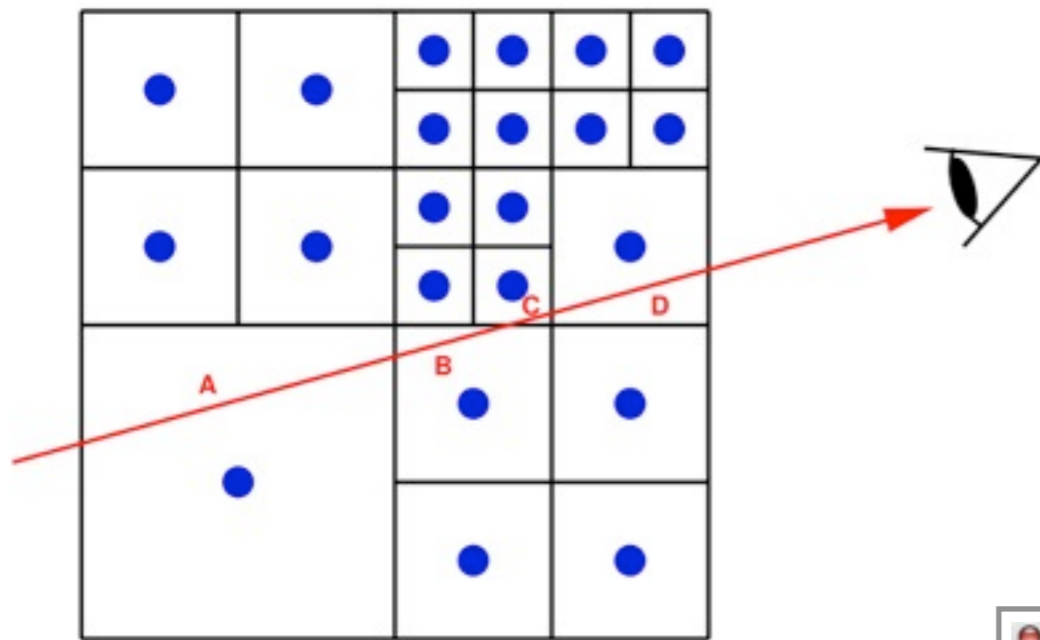


Lyutikov et al. (2005)

The RADMC-3D Code:



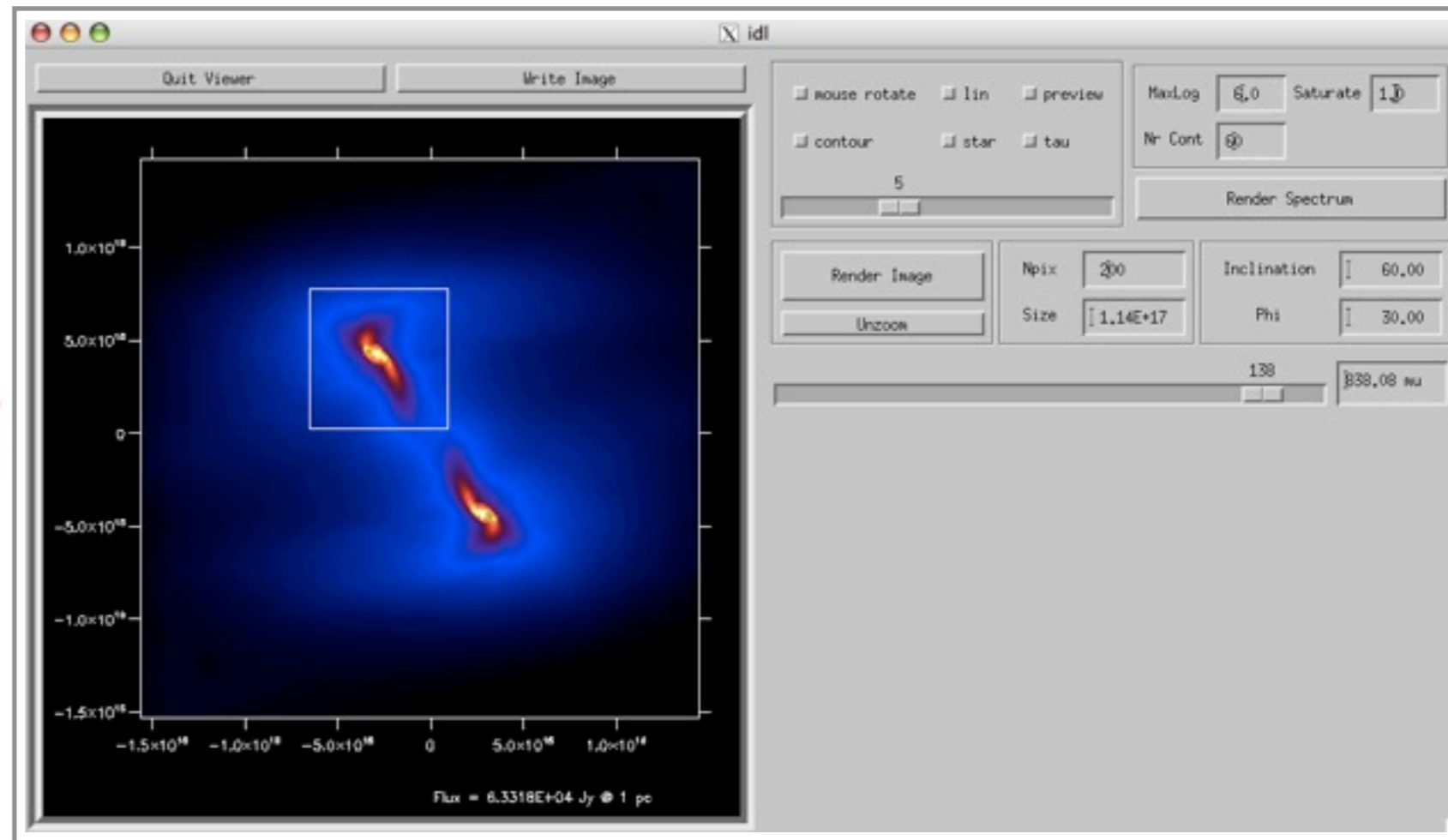
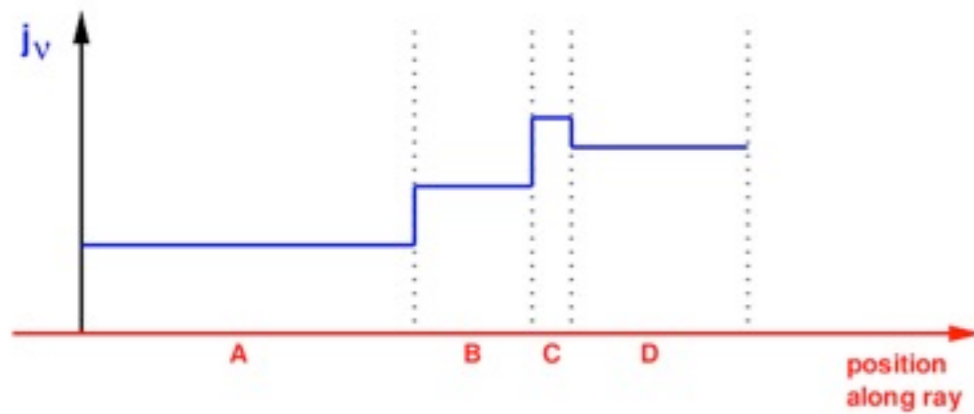
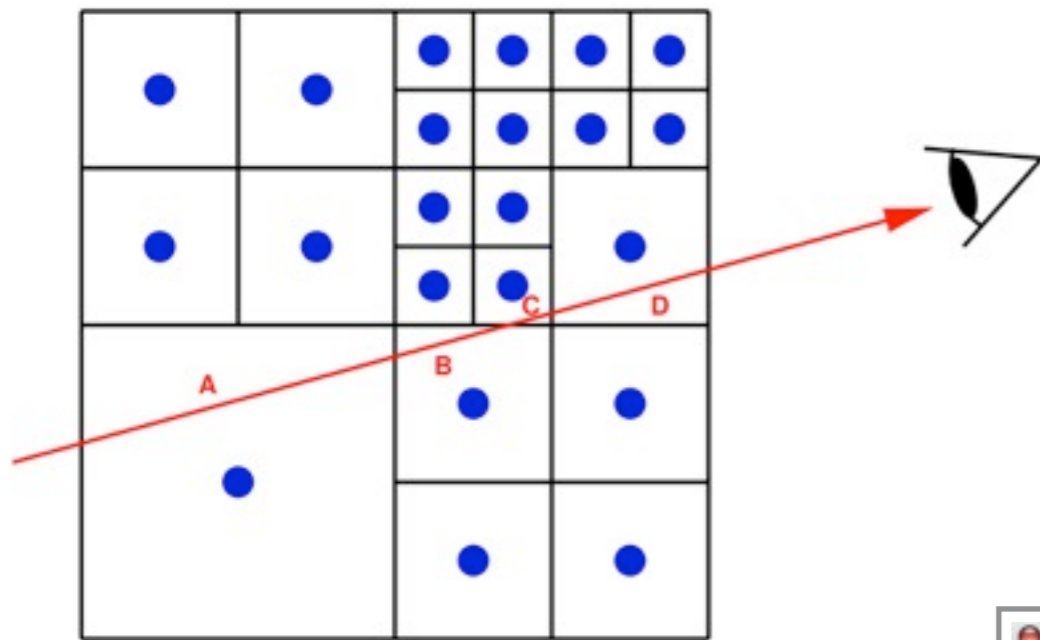
MAX-PLANCK-GESELLSCHAFT



The RADMC-3D Code:



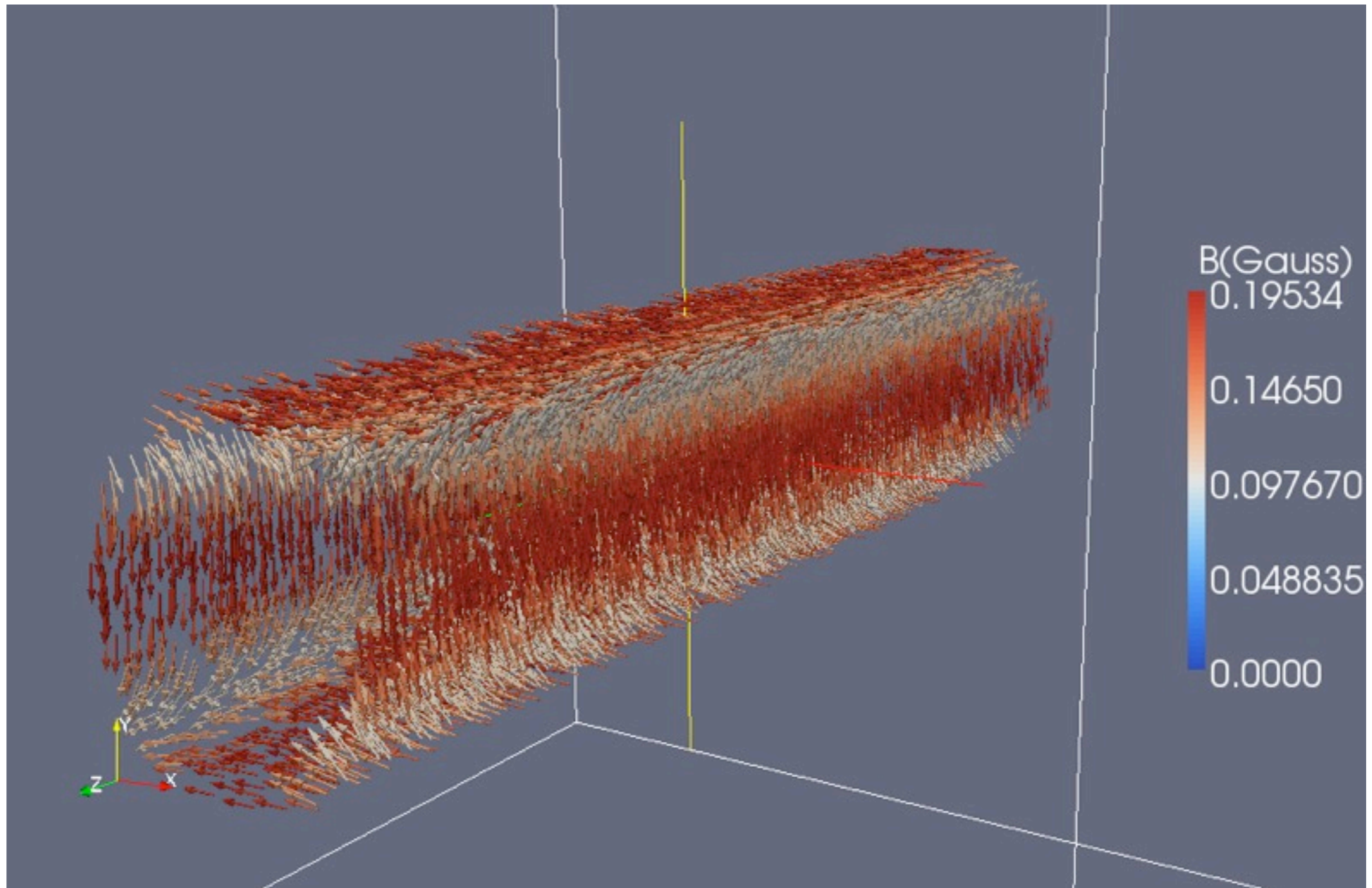
MAX-PLANCK-GESELLSCHAFT



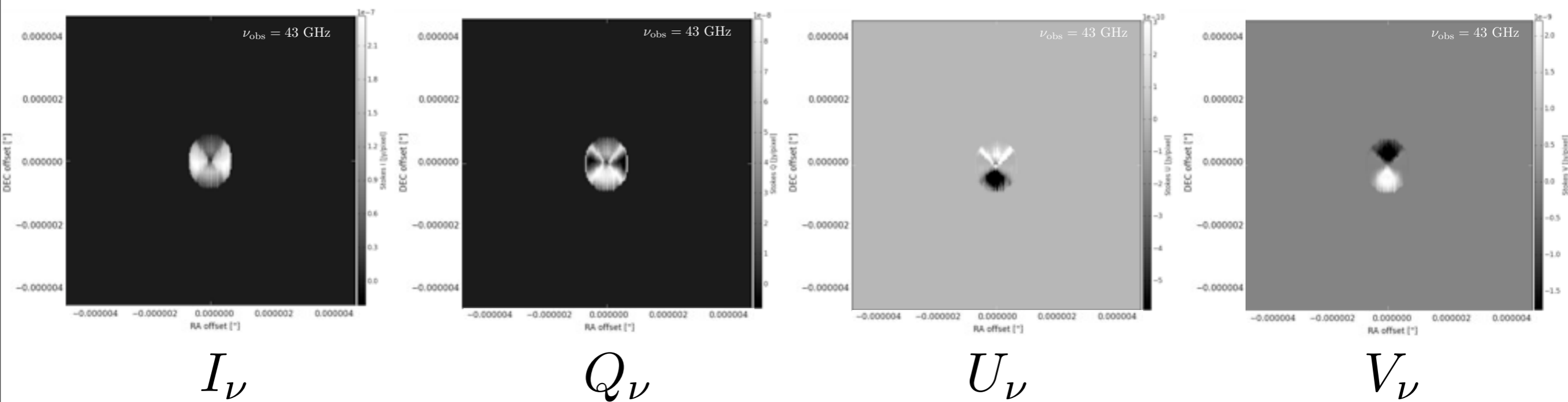
radmc3dPy

Attila Juhasz

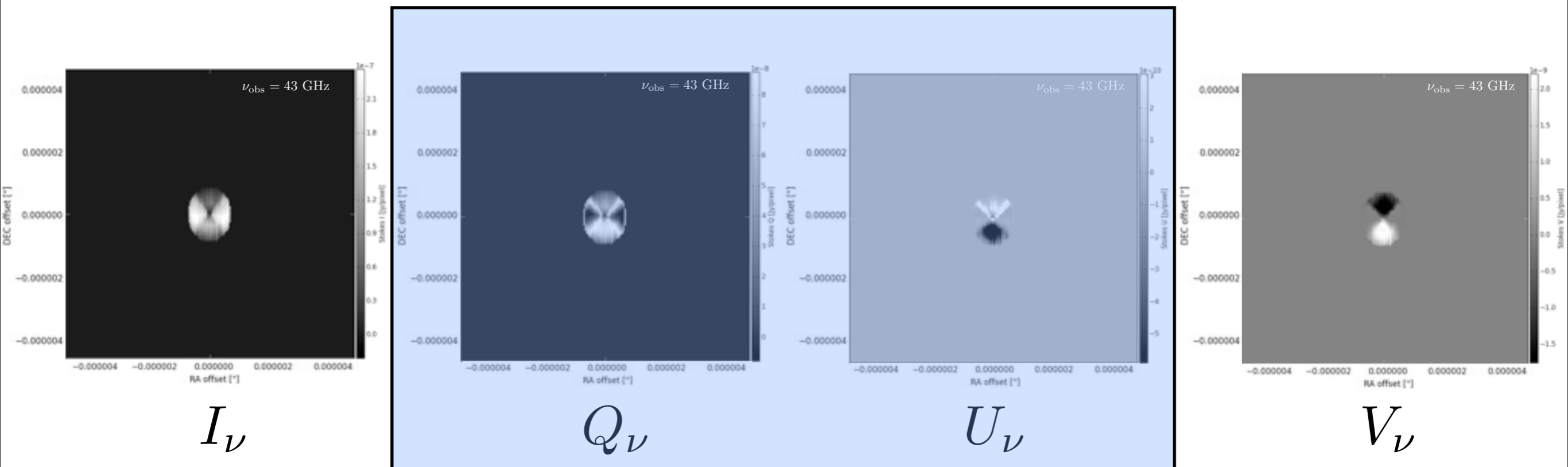
TEMZ Model (Order):



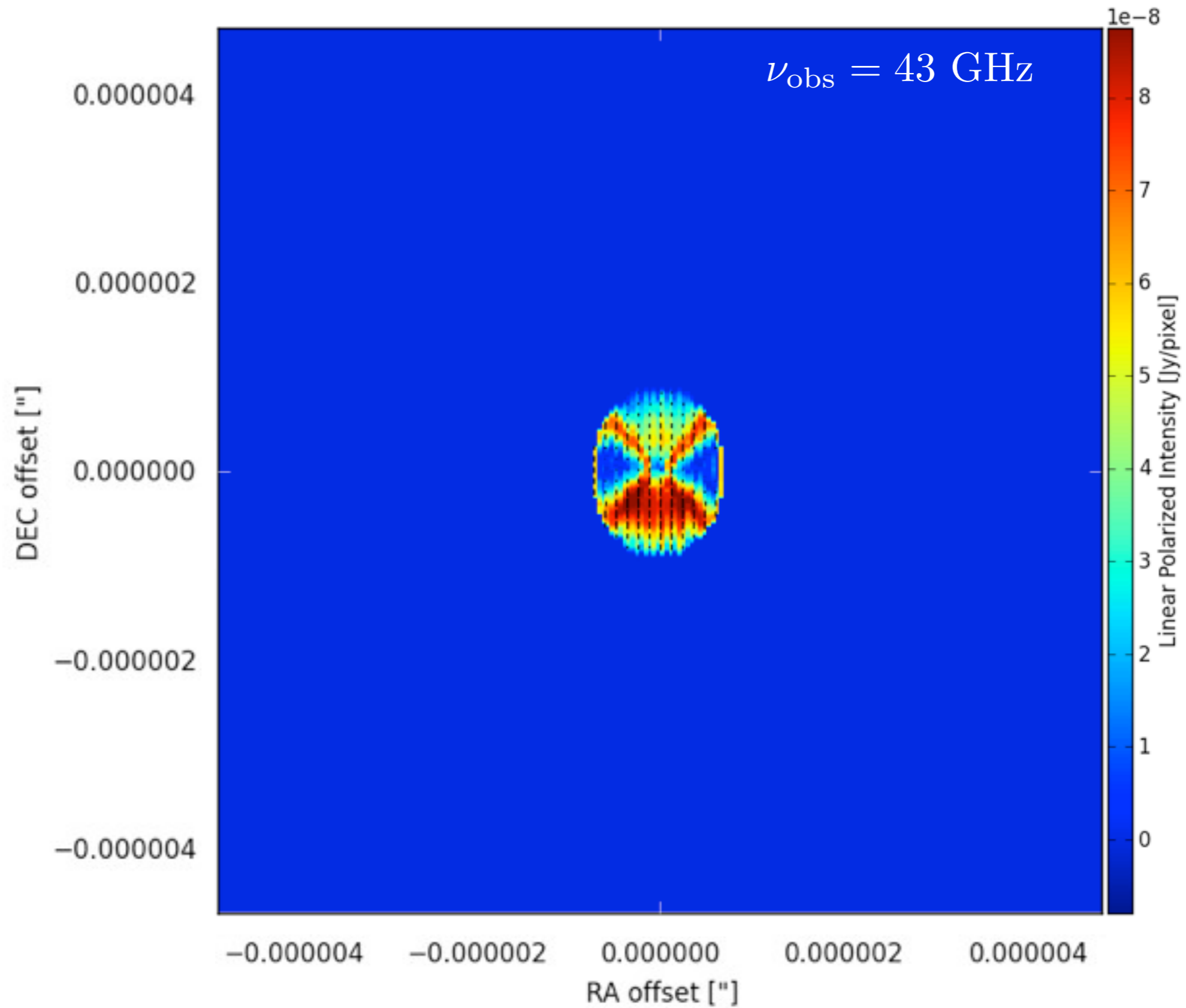
TEMZ Model (Order):



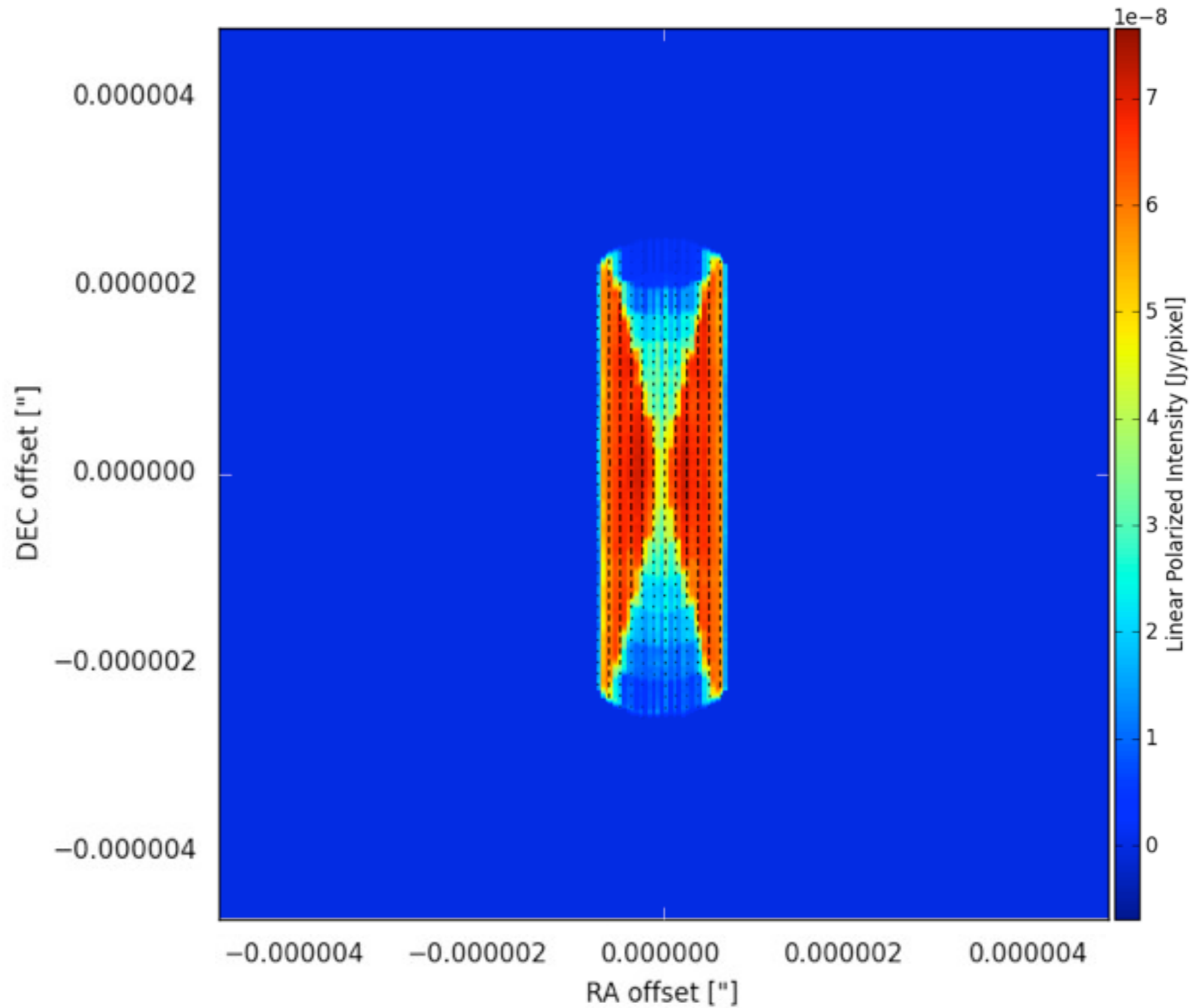
TEMZ Model (Order):



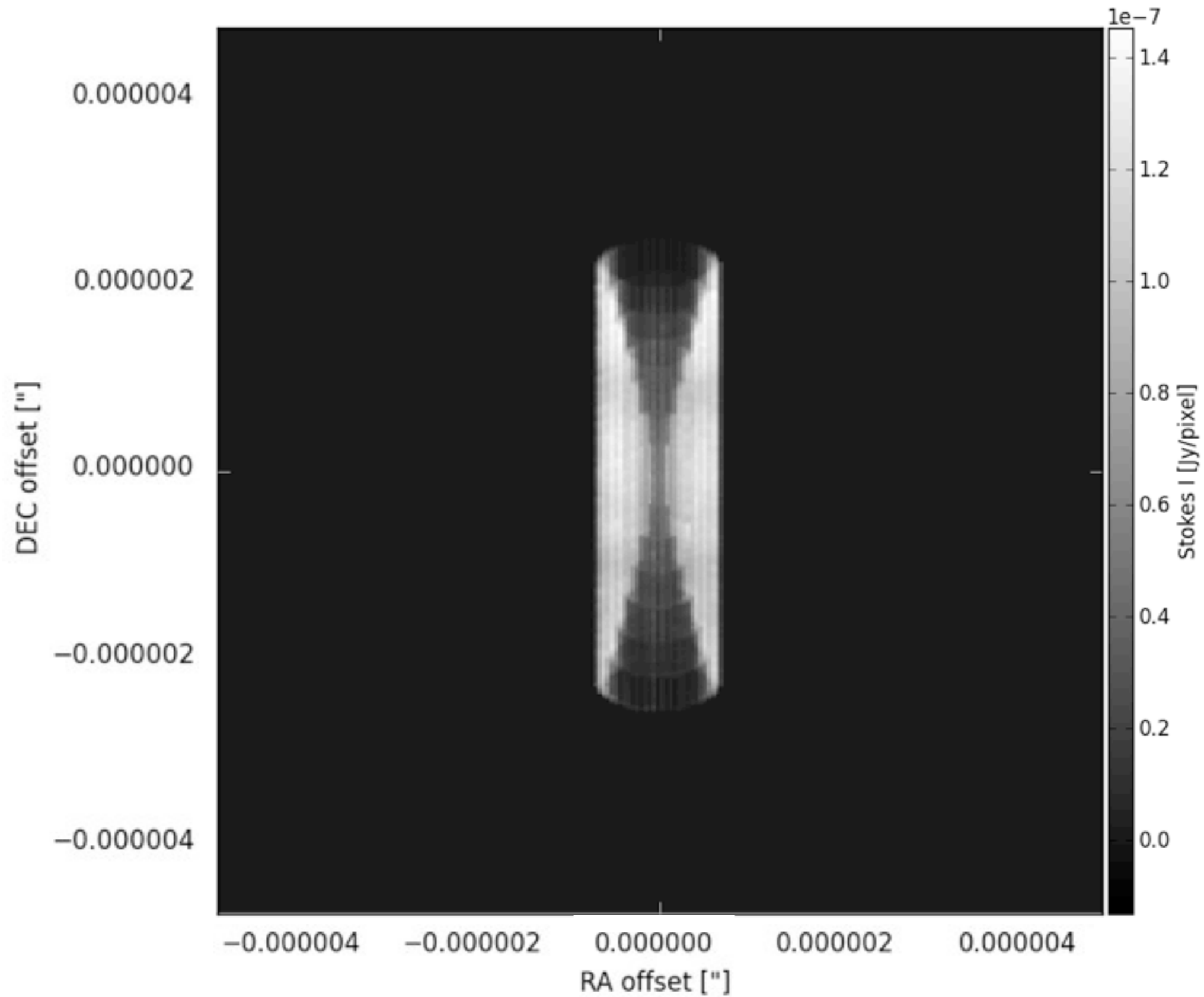
TEMZ Model (Order):



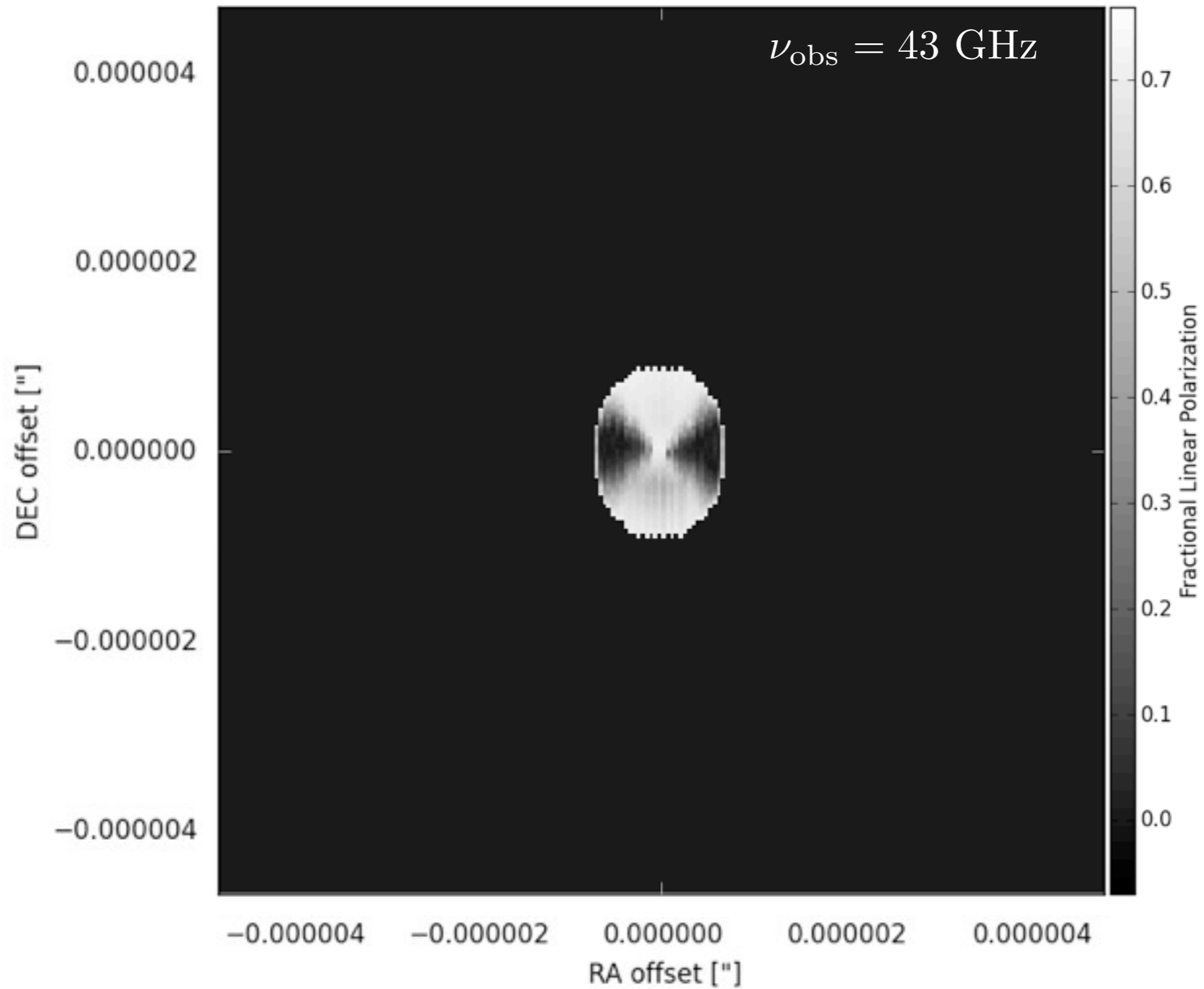
TEMZ Model (Order):



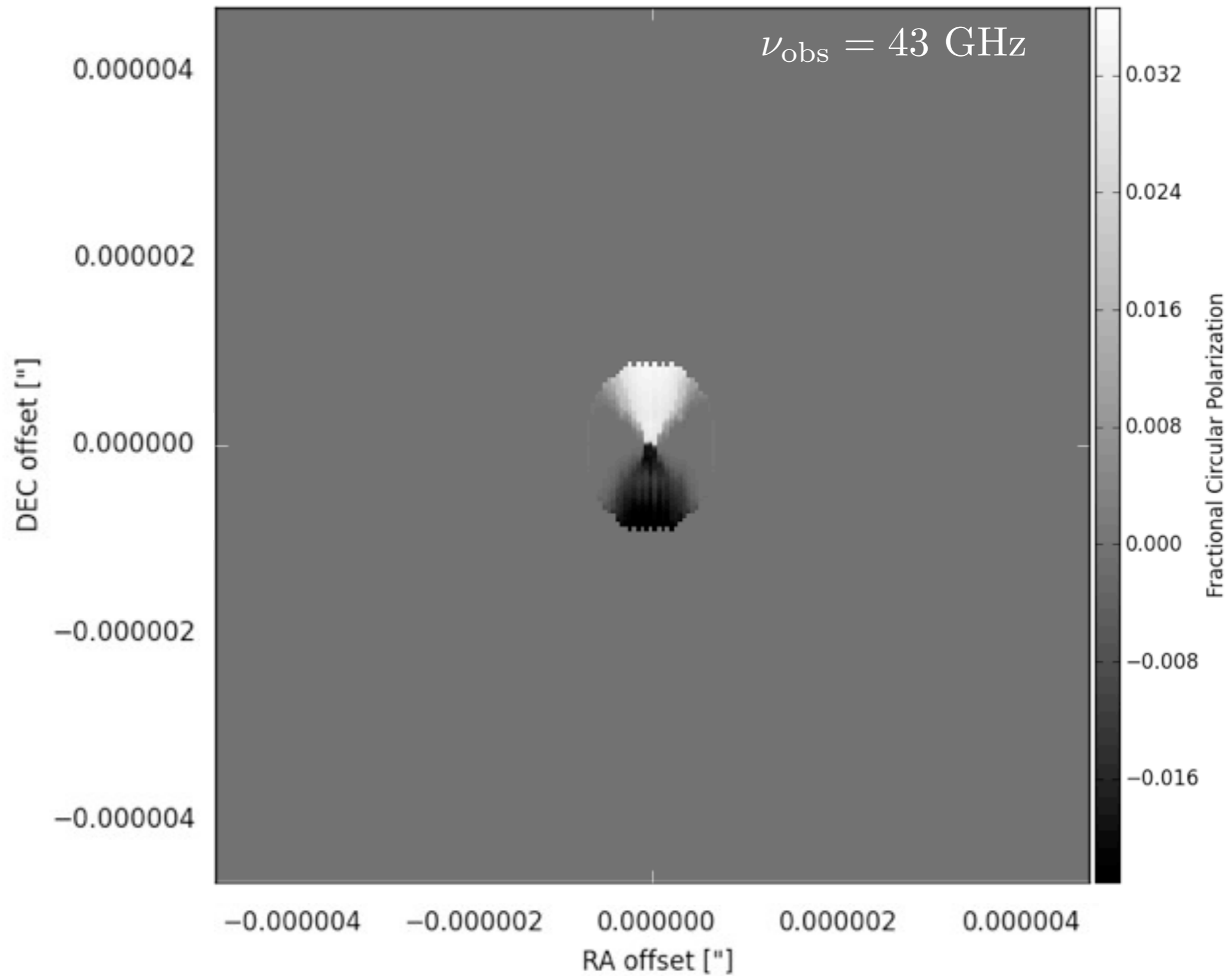
TEMZ Model (Order):



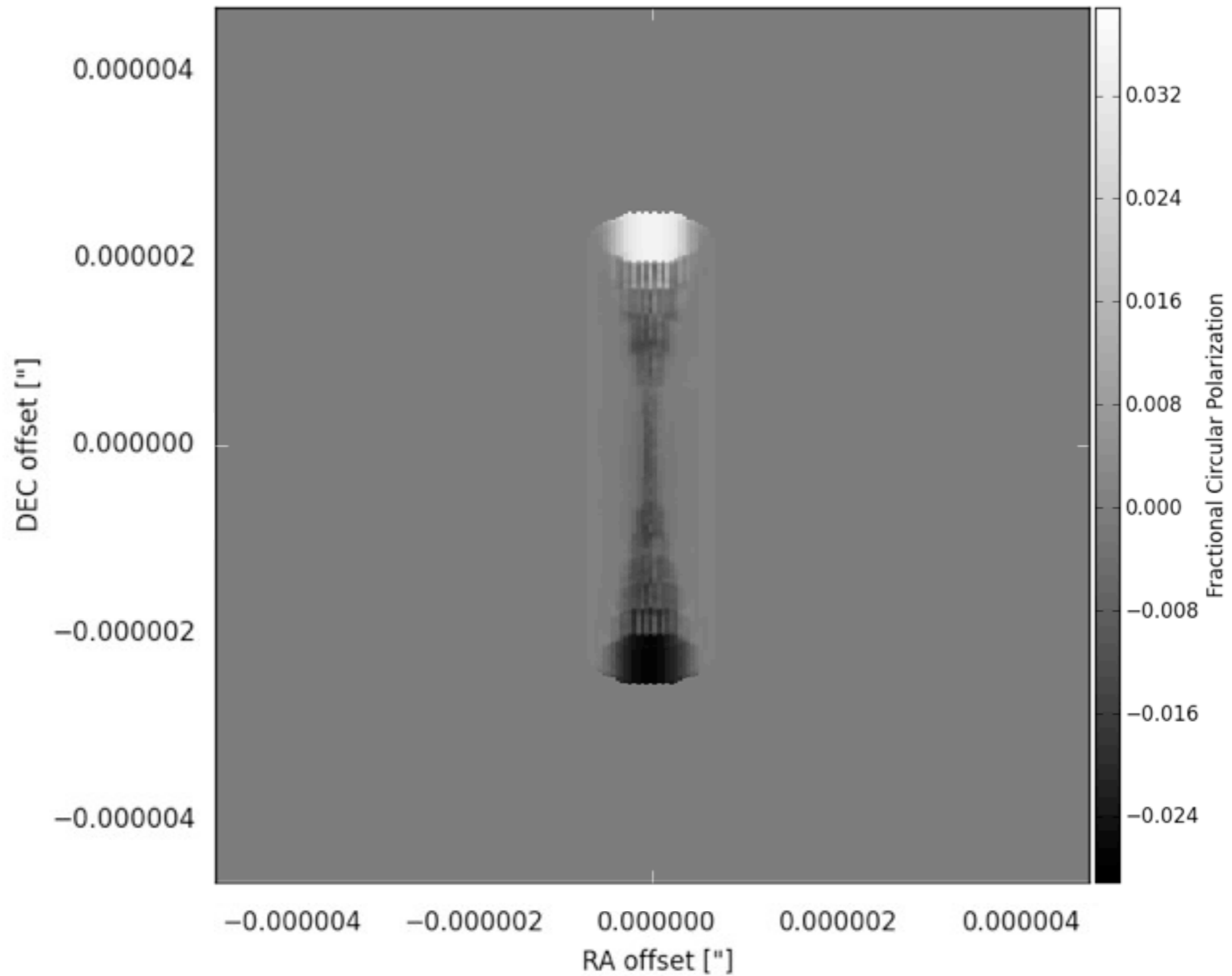
TEMZ Model (Order):



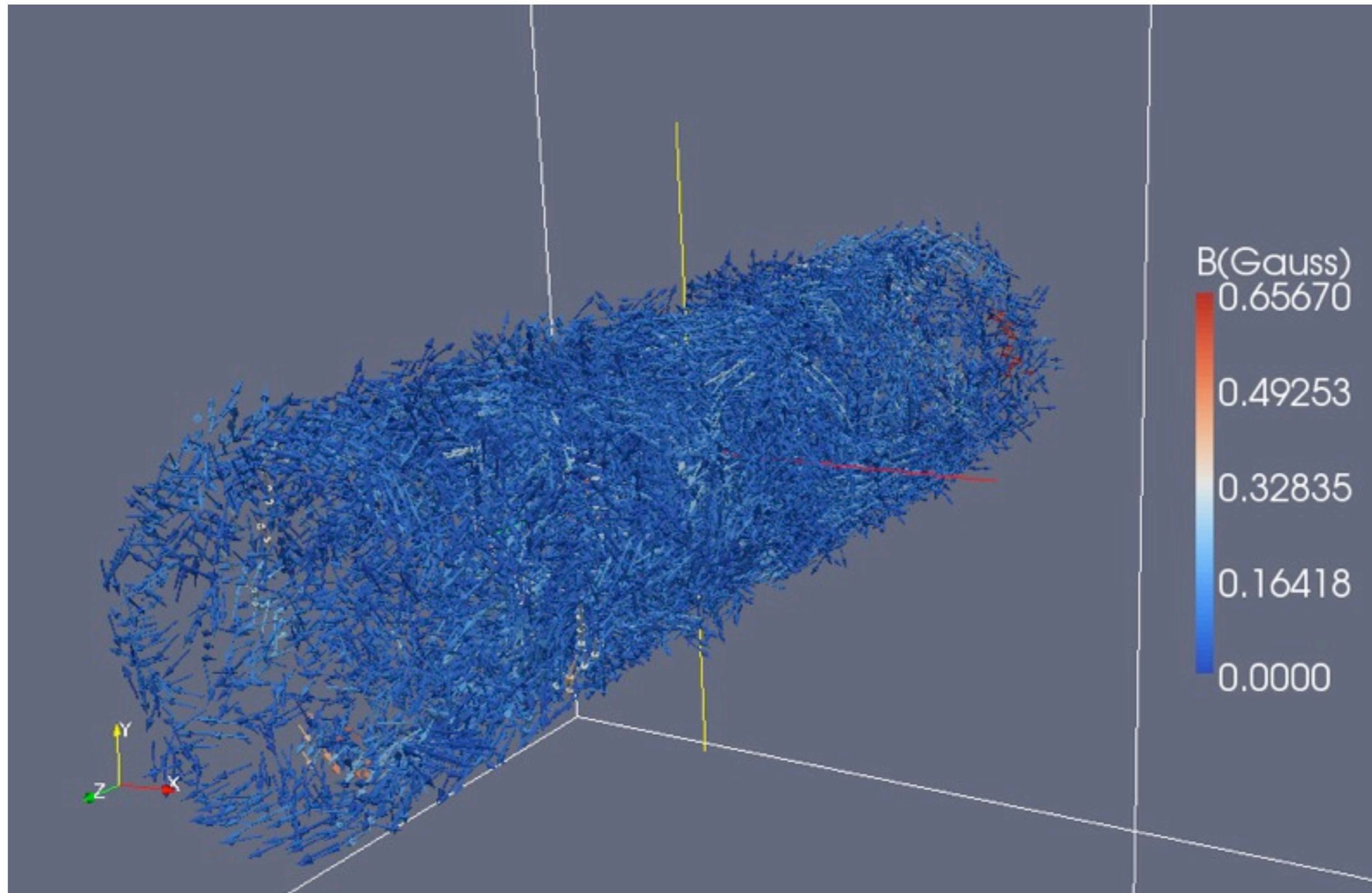
TEMZ Model (Order):



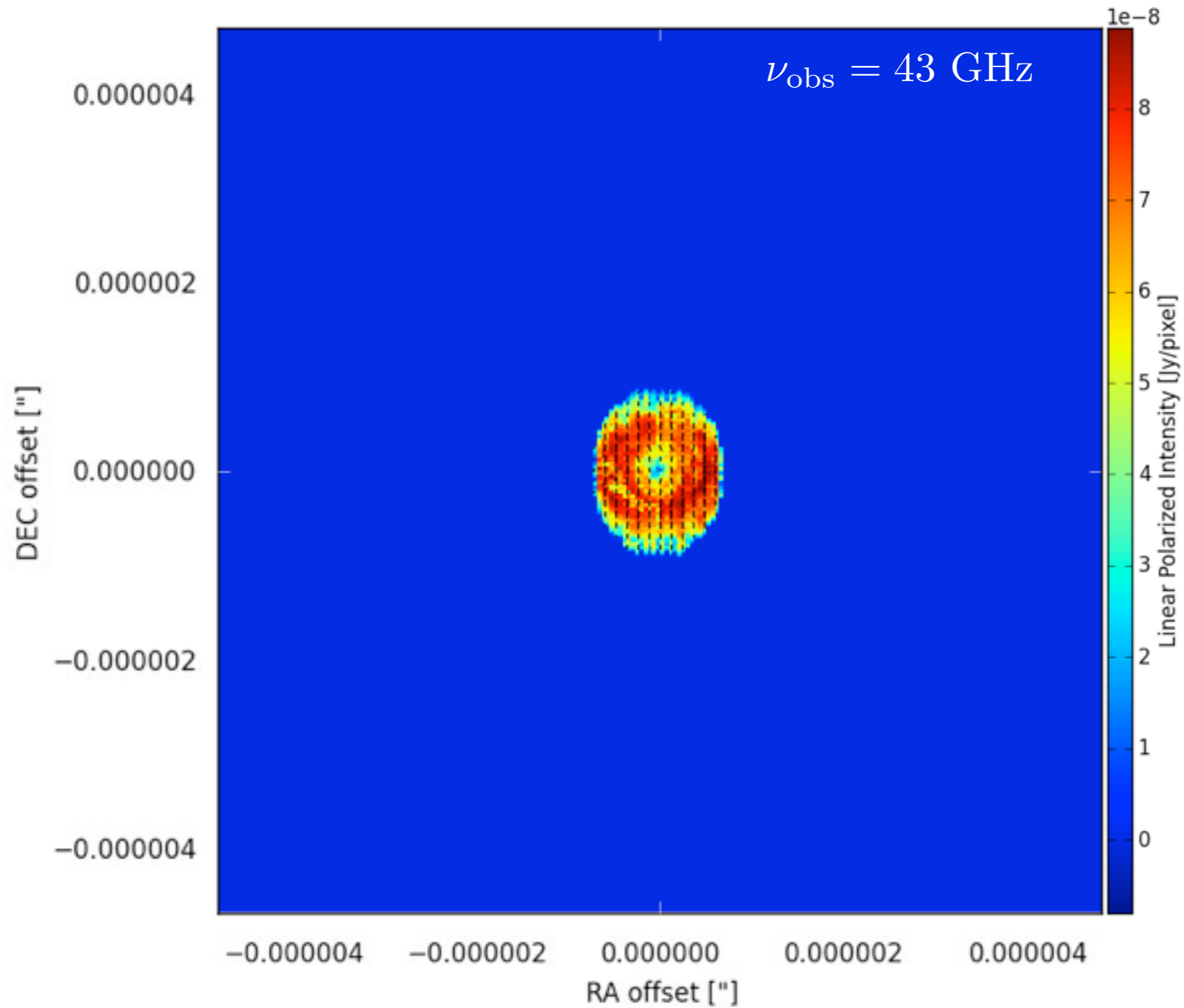
TEMZ Model (Order):



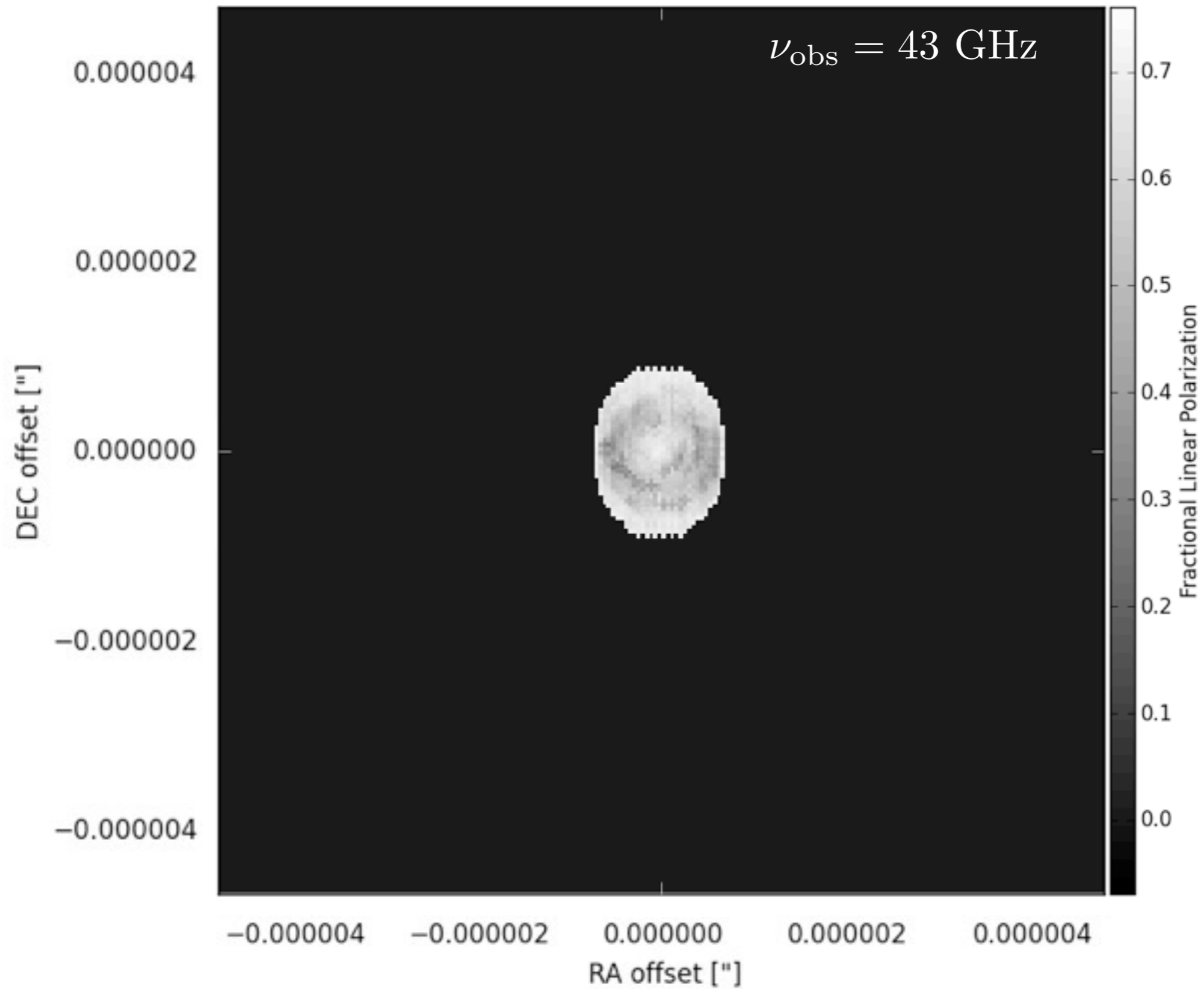
TEMZ Model (Disorder):



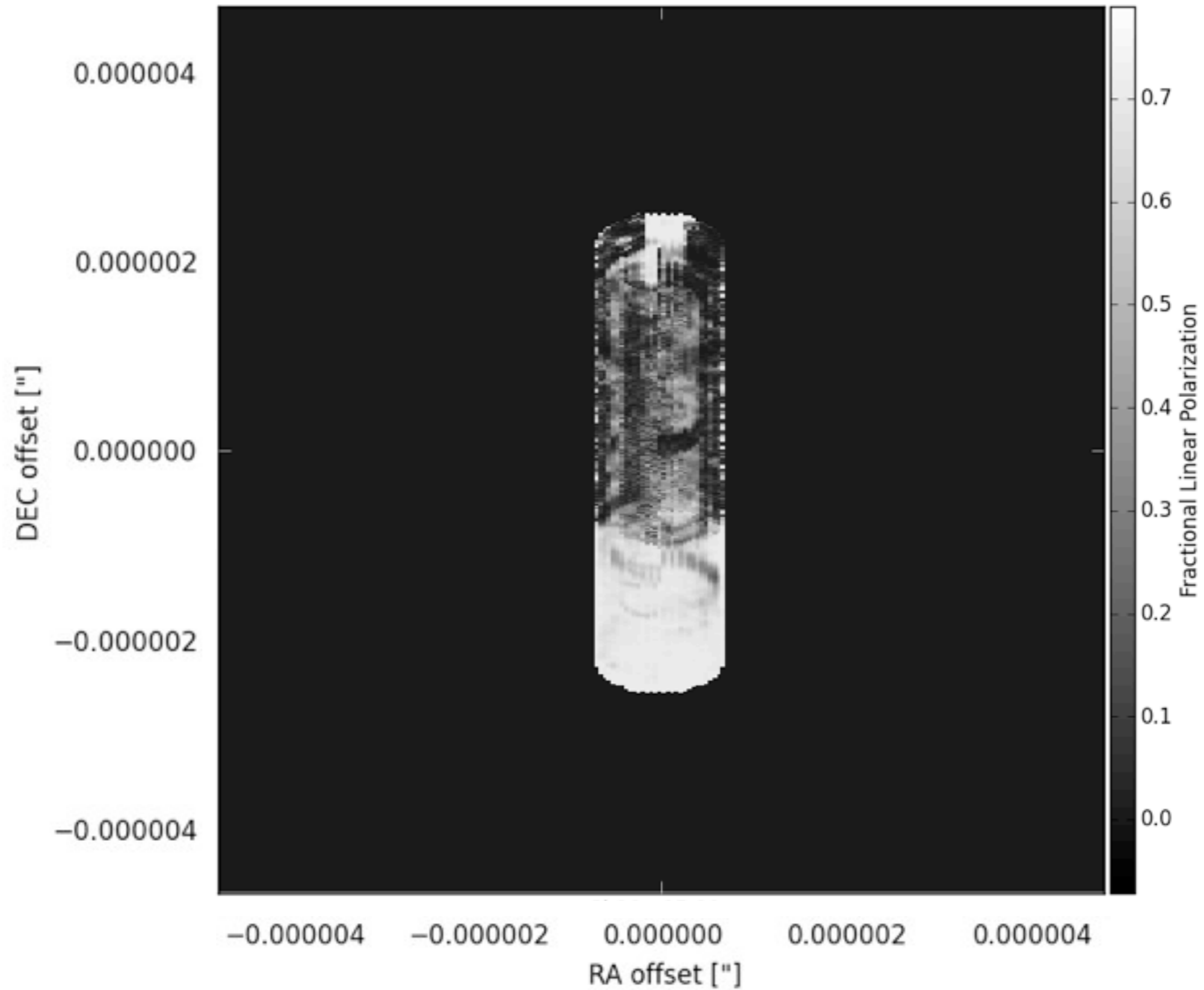
TEMZ Model (Disorder):



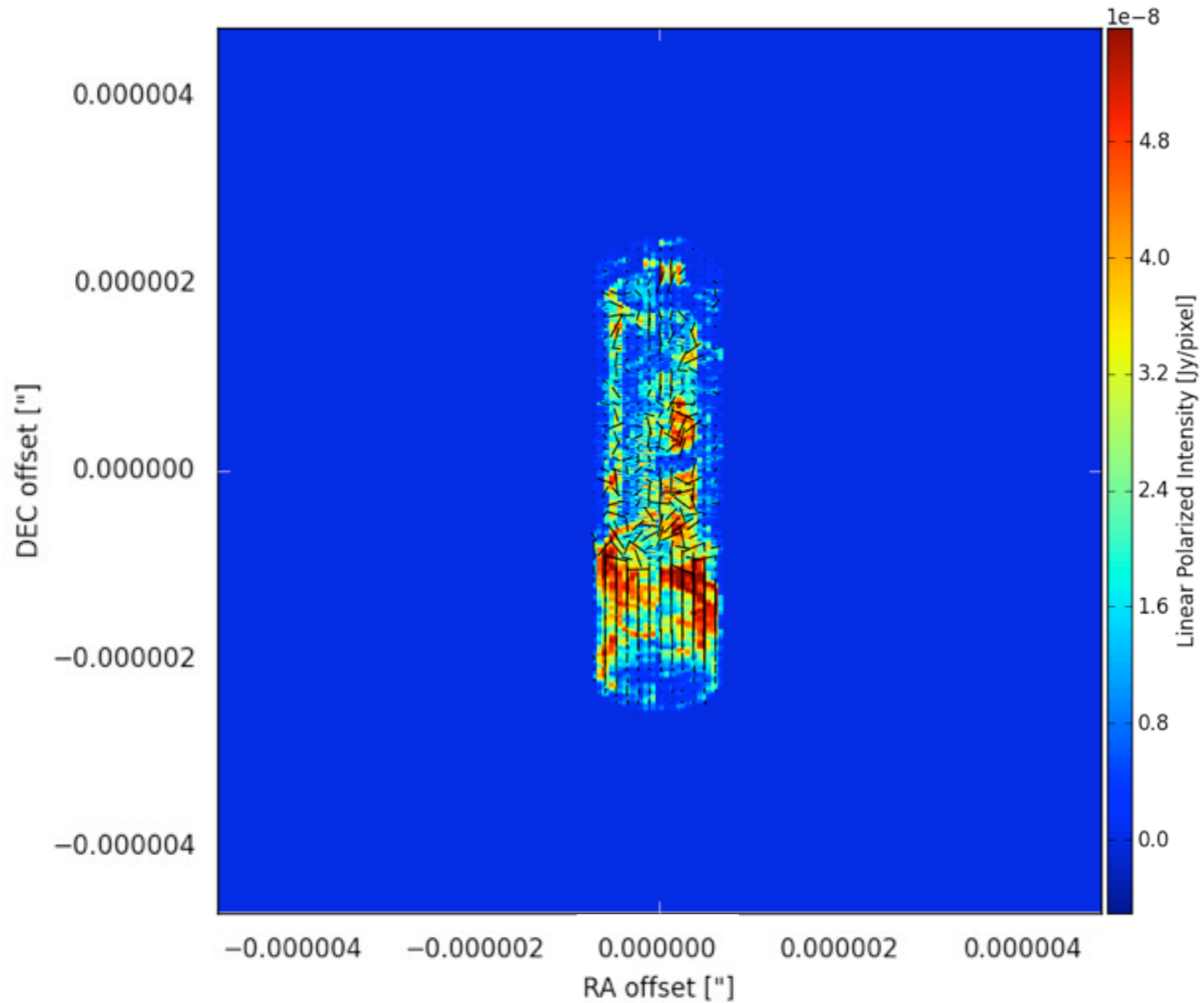
TEMZ Model (Disorder):



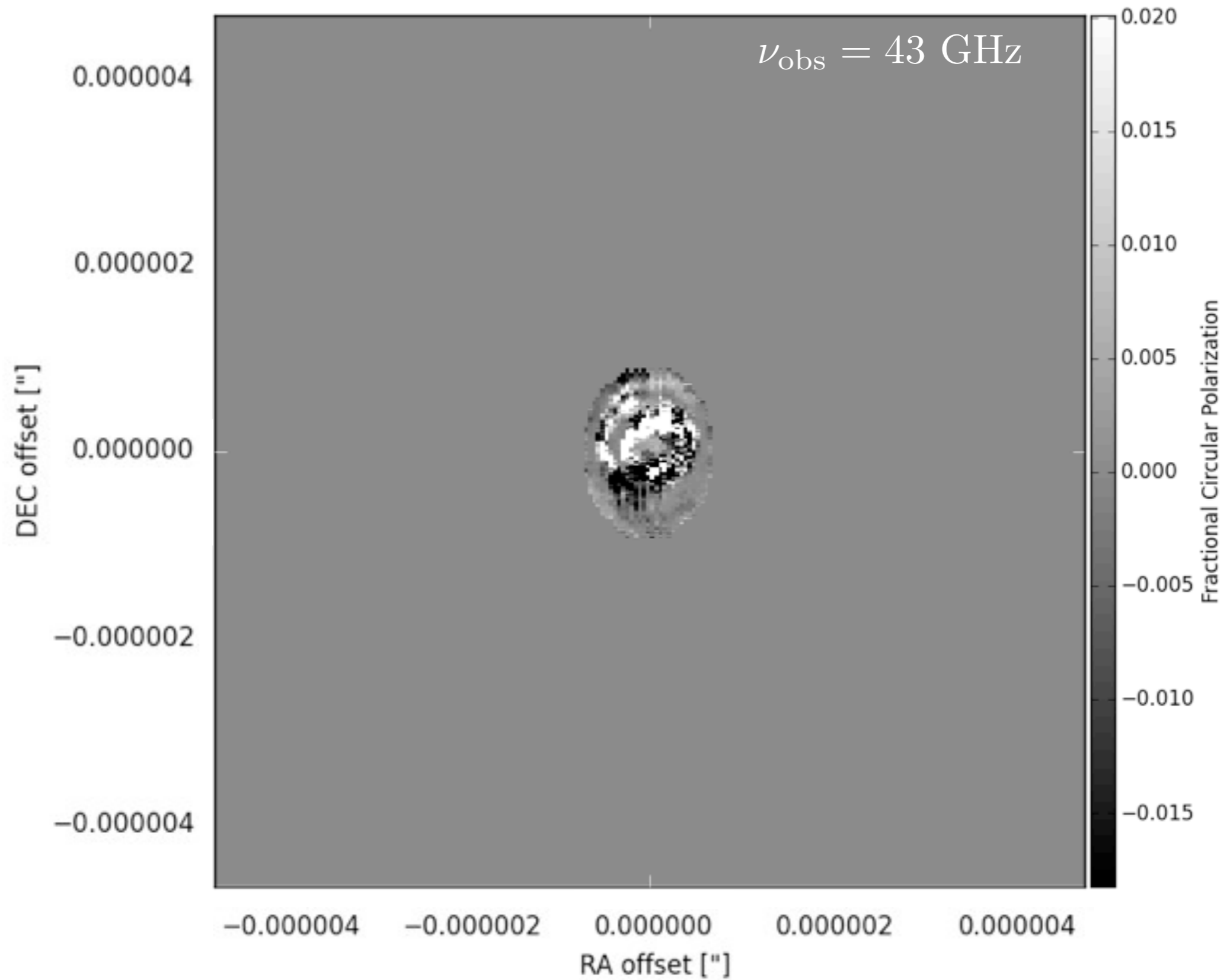
TEMZ Model (Disorder):



TEMZ Model (Disorder):



TEMZ Model (Disorder):



Can circular polarization be produced within a turbulent jet?

A large, swirling galaxy in shades of blue and cyan, set against a dark blue background filled with stars. The galaxy's core is bright and emits a thin, white beam of light that extends downwards. The word "Questions?" is written in a white, sans-serif font in the center of the galaxy's core.

Questions?